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Cosmic Dawn (CoDa): the First Radiation-Hydrodynamics Simulation of Reionization and Galaxy Formation in the Local Universe

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ABSTRACT

Cosmic reionization by starlight from early galaxies affected their evolution, thereby impacting reionization, itself. Star formation suppression, for example, may explain the observed underabundance of Local Group dwarfs relative to N-body predictions for Cold Dark Matter. Reionization modelling requires simulating volumes large enough [$\sim (100 \text{ Mpc})^3$] to sample reionization “patchiness”, while resolving millions of galaxy sources above $\sim 10^8 M_\odot$, combining gravitational and gas dynamics with radiative transfer. Modelling the Local Group requires initial cosmological density fluctuations pre-selected to form the well-known structures of the local universe today. Cosmic Dawn (“CoDa”) is the first such fully-coupled, radiation-hydrodynamics simulation of reionization of the local universe. Our new hybrid CPU-GPU code, RAMSES-CUDATON, performs hundreds of radiative transfer and ionization rate-solver timesteps on the GPUs for each hydro-gravity timestep on the CPUs. CoDa simulated $(91 \text{ Mpc})^3$ with 4096^3 particles and cells, to redshift 4.23, on ORNL supercomputer Titan, utilizing 8192 cores and 8192 GPUs. Global reionization ended slightly later than observed. However, a simple temporal rescaling which brings the evolution of ionized fraction into agreement with observations also reconciles ionizing flux density, cosmic star formation history, CMB electron scattering optical depth and galaxy UV luminosity function with their observed values. Photoionization heating suppressed the star formation of haloes below $\sim 2 \times 10^9 M_\odot$, decreasing the abundance of faint galaxies around $M_{AB1600} = [-10, -12]$. For most of reionization, star formation was dominated by haloes between $10^{10} - 10^{11} M_\odot$, so low-mass halo suppression was not reflected by a distinct feature in the global star formation history. Intergalactic filaments display sheathed structures, with hot envelopes surrounding cooler cores, but do not self-shield, unlike regions denser than $100 \langle \rho \rangle$.

1 INTRODUCTION

The epoch of reionization (hereafter, “EoR”) resulted from the escape of ionizing radiation from the first generation of star-forming galaxies into the otherwise cold, neutral gas of the primordial intergalactic medium (“IGM”). This radiation created intergalactic H II regions of ever-increasing size, surrounding the galaxies that created them, leading eventually to the complete overlap of neighboring H II regions and the end of the EoR within the first billion years after the big bang. The growth and geometry of the H II regions reflected that of the underlying galaxies as the radiation sources and their spatial clustering, and that of the density fluctuations in the intergalactic gas as the absorbing medium. The story of cosmic reionization is inseparable, therefore, from that of the emergence of galaxies and large-scale structure in the universe. Recently, the observational frontier has begun to push the look-back-time horizon accessible to direct observation back into this period. High-redshift galaxies and quasars have been observed, for example, which constrain the evolution of the IGM opacity to H Lyman alpha resonance scattering, the mean ionizing flux density, the UV luminosity function of galaxies, and the cosmic star formation history, while observations of the cosmic microwave background (“CMB”) constrain the mean electron scattering optical depth integrated thru the entire IGM over time, back to the epoch of recombination, and give some information about the evolution of the mean ionized fraction. Observational probes of the large-scale structure of reionization and its evolving “patchiness” are just beginning. The cosmic 21cm background from the evolving patchwork of intergalactic H I regions, not yet reionized, is a prime example, which is an important science-driver for the development of a new generation of low-frequency radio telescope arrays, such as LOFAR, MWA, PAPER, and SKA.

The EoR is also thought to be important for its impact on galaxy formation. It has been suggested that the rising intergalactic UV radiation field is responsible for suppressing the star formation of low-mass galaxies, by removing gas from the smallest galaxies — minihalos — by photoevaporation (Shapiro et al. 2004; Iliev et al. 2005) and through the suppression of gas infall onto low-mass galaxies above the minihalo mass range, by photoheating the intergalactic gas and raising its pressure (Shapiro et al. 1994; Gnedin 2000; Hoesft et al. 2006), affecting their star formation efficiency.

This process could provide a credible solution to the “missing satellites problem” (Klypin et al. 1999; Moore et al. 1999), by inhibiting star formation in low mass galaxies at early times (Bullock et al. 2000; Benson et al. 2002b,a, 2003). In this framework, a number of semi-analytical models (hereafter SAMs) have been shown to reproduce well the satellite population of the Milky Way (hereafter MW), such as Koposov et al. (2009); Muñoz et al. (2009); Busha et al. (2010); Macciò et al. (2010); Li et al. (2010); Font et al. (2011). They suggest that the ultra-faint dwarf galaxies (hereafter UFDs) discovered by the SDSS (Martin et al. 2004; Willman & et al. 2005; Zucker et al. 2006; Belokurov & et al. 2007; Irwin & et al. 2007; Walsh et al. 2007) are effectively reionisation fossils, living in dark matter sub-haloes of about $10^{6-9}M_{\odot}$. Deep HST observations have recently made it possible to start testing this idea. While this is currently a matter of

debate, it seems that at least the low-mass satellites of the MW and M31 may have star formation histories compatible with a very early suppression (Brown et al. 2014; Weisz et al. 2015), e.g. by reionization. In this context, we are led to the exciting possibility that low-mass satellites in the Local Group can be used as probes of the local reionization process.

While the reionization of the Universe is often considered to be complete at $z = 6$ (Fan et al. 2006), further quasar Lyman α observations have shown that extended opaque hydrogen troughs may still exist below this redshift, highlighting the patchiness of the process and suggesting that the reionization may actually still be ongoing at $z = 6$ and end as low as $z \sim 5$ (Becker et al. 2015). Given this observational evidence, one can only expect that the reionization of the Local Group must have been very different from a uniform or instantaneous event. Indeed, Iliev et al. (2011) described two distinct reionization scenarios for the Local Group, internal (and slow, driven by sources within the Local Group) or external and fast, driven by a massive, rare object such as Virgo. Many more scenarios can be conceived, depending on the physics considered (e.g. quasars, X-ray binaries), and it would be desirable, as a long term goal, to establish which properties of the local satellites population could allow us to discriminate between these scenarios.

Already, Ocvirk & Aubert (2011) showed that the structure of the UV background during reionization has a strong impact on the properties of the satellite population of galaxies. In particular, they showed that an internally-driven reionization led to significant changes in the radial distribution of the satellites of the MW. This prediction was then further supported by using more refined models (Ocvirk et al. 2013, 2014), built upon high resolution zoom simulations of the formation of the Local Group performed by the CLUES project¹. However, these studies suffered from several limitations. First of all, these studies were performed in the so-called “post-processing” framework, and therefore only crudely account for stellar feedback (radiative and SNe). Furthermore, since the gas density field is frozen in the post-processing paradigm, the effect of radiation on the distribution of gas in the IGM (smoothing of gas overdensities due to photo-heating) can not be accounted for either, while it has been shown to have a significant impact on photon consumption (Pawlik et al. 2015). Second, the gas distribution is only known in the zoom region, which is only a few Mpc wide, and mainly contains the progenitors of the MW, M31 and M33, and the IGM in-between. Beyond this region, the box is populated with low resolution pure dark matter particles. This dual description makes it very challenging to describe external (i.e. pure dark matter, low resolution) and internal (i.e. zoom region, including gas particles) reionization sources in a homogeneous, consistent way.

Our ultimate challenge for modelling reionization, then, is to be able to simulate the coupled multi-scale problem of global reionization and individual galaxy formation, with gas and gravitational dynamics and radiative transfer, simultaneously. And to predict the relic impact of this global reion-

¹ Constrained Local UniversE Simulations, <http://www.clues-project.org/index.html>

ization on the universe today, to make comparisons possible with the observationally-accessible nearby universe, including the Local Group and its satellites, it is further necessary to start from initial conditions preselected to produce the observed galaxies and large-scale structure of the local universe. For this simulation to characterize the evolving "patchiness" of reionization in a statistically meaningful way requires a comoving simulation volume as large as $\sim (100 \text{ Mpc})^3$ [e.g. (Iliev et al. 2006, 2014)]. Accounting for the millions of galaxies in this volume over the full mass range of galactic haloes which are thought to contribute most significantly to reionization, the so-called "atomic-cooling haloes" (henceforth, "ACHs"), those with virial temperatures above $\sim 10^4 \text{ K}$ (corresponding to halo masses above $\sim 10^8 \text{ M}_\odot$) while modelling the impact of reionization on these individual galaxies and the IGM, requires a physical resolution of a few kpc's over the entire volume.

To meet this challenge and satisfy all these requirements, we have developed a new, hybrid CPU-GPU, fully-coupled, cosmological radiation-hydrodynamics-gravity code, RAMSES-CUDATON, capable of simulating the EOR in a comoving volume 91 Mpc on a side, with 4096^3 dark matter particles and a cubic-lattice of 4096^3 cells for the gas and radiation field. Our simulation self-consistently models BOTH global reionization AND the formation and reionization of the Local Group, by starting from a "constrained realization" of initial conditions for the local universe which reproduces all the familiar features of the local universe in a volume centered on the Local Group, such as the Milky Way, M31, and the Fornax and Virgo clusters — i.e. CLUES initial conditions. We call this simulation "CoDa", for Cosmic Dawn.

The main goal of this first paper is to introduce the simulation and compare the results with current observational constraints on the EOR, such as the evolution of the cosmic neutral fraction, the cosmic ionizing flux density, the cosmic star formation rate, the Thomson optical depth measured by cosmic microwave background experiments and high redshift UV luminosity functions (Sec. 3). It is laid out as follows: first we describe the code and the simulation setup (Sec. 2). We then proceed to our results and compare to available observational constraints of the EoR (Sec. 3), and finish with a short summary.

2 METHODOLOGY

The Cosmic Dawn simulation uses the fully coupled radiation hydrodynamics code RAMSES-CUDATON. This section describes the principles of the code and its deployment. For quick reference, the parameters of the simulation are summarized in Tab. 1.

2.1 RAMSES-CUDATON

2.1.1 RAMSES

RAMSES is a code for simulating the formation of large scale structures, galaxy formation and self-gravitating hydrodynamics in general (Teyssier 2002). Collisionless N-body dynamics are solved via a particle-mesh integrator. The adaptive mesh refinement was turned off in this run (see Sec. 2.1.3

| Cosmology (WMAP5+BAO+SN) | |
|---|---|
| Dark energy density Ω_Λ | 0.721 |
| Matter density Ω_m | 0.279 |
| Baryonic matter density Ω_b | 0.046 |
| Hubble constant $h = H_0/(100 \text{ km/s})$ | 0.70 |
| Power spectrum | |
| Normalization σ_8 | 0.817 |
| Index n | 0.96 |
| Setup | |
| Number of nodes (GPUs) | 8192 (8192) |
| Grid size | 4096^3 |
| Box size L_{box} | 91 Mpc |
| Grid cells per node | $128 \times 256 \times 256$ |
| DM particle number N_{DM} | 4096^3 |
| DM particle mass M_{DM} | $3.49 \times 10^5 \text{ M}_\odot$ |
| Initial redshift z_{start} | 150 |
| End redshift z_{end} | 4.23 |
| Star formation | |
| Density threshold δ_\star | 50 $\langle \rho \rangle$ |
| Temperature threshold T_\star | $2 \times 10^4 \text{ K}$ |
| Efficiency ϵ_\star | 10^{-2} |
| Stellar particle mass (post-SN) M_\star | 3194 M_\odot |
| Massive star lifetime t_\star | 10 Myr |
| Supernova feedback | |
| Mass fraction η_{SN} | 10% |
| Energy E_{SN} | 10^{51} erg |
| Radiation | |
| Stellar ionizing emissivity | $1.824 \times 10^{47} \text{ ph/s/M}_\odot$ |
| Stellar particle escape fraction $f_{\text{esc},\star}$ | 0.5 |
| Effective photon energy | 29.61 eV |
| Effective H cross-section σ_E | $1.097 \times 10^{-22} \text{ m}^2$ |
| Speed of light c | 299 792 458 m/s |

Table 1. Cosmic Dawn simulation parameters summary

for more explanations on this). Gas dynamics is modeled using a second-order unsplit Godunov scheme (Teyssier et al. 2006; Fromang et al. 2006) based on the HLLC Riemann solver (Toro et al. 1994). We assume a perfect gas Equation of State (hereafter EoS) with $\gamma = 5/3$.

We consider star formation using a phenomenological approach. In each cell with gas density larger than a gas overdensity $\delta_\star = 50$, we spawn new star particles at a rate given by

$$\dot{\rho}_\star = \epsilon_\star \frac{\rho_{gas}}{t_{ff}} \text{ with } t_{ff} = \sqrt{\frac{3\pi}{32G\rho}} \quad (1)$$

where t_{ff} is the free-fall time of the gaseous component and $\epsilon_\star = 0.01$ is the star formation efficiency.

We also require the cell temperature to be lower than $T_\star = 2 \times 10^4 \text{ K}$ in order to form stars: above this temper-

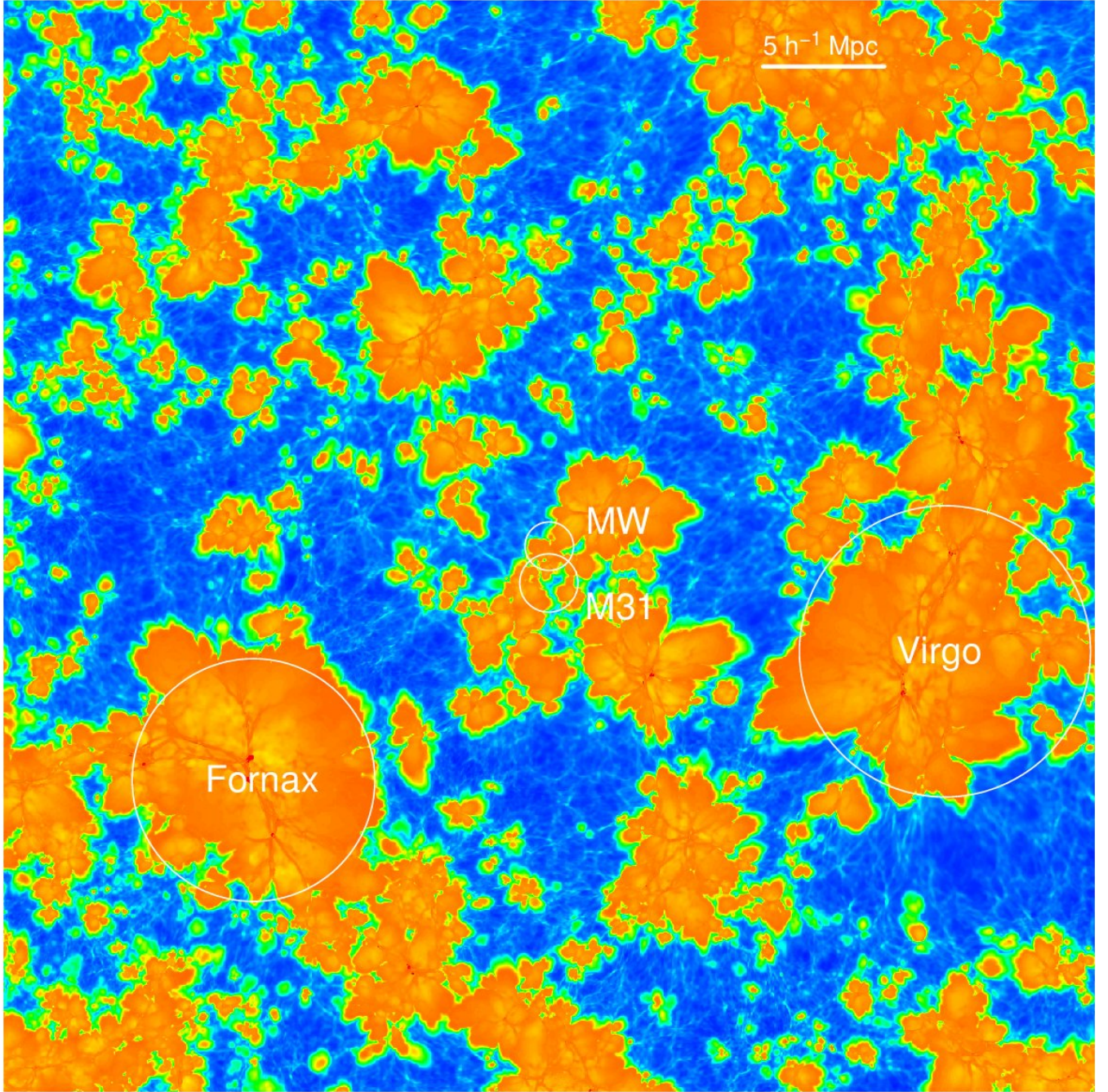


Figure 1. Temperature distribution in a $45 \, h^{-1} \text{Mpc} \times 45 \, h^{-1} \text{Mpc} \times 0.03 \, h^{-1} \text{Mpc}$ slice of the simulation at redshift $z_{\text{sim}} = 6.15$. Orange regions show photo-heated, ionized material, while the cold, still neutral medium appears in blue. The slice is situated in the supergalactic YZ plane with its origin shifted to the center on the Local Group progenitor at $z_{\text{sim}} = 7$. The circles of decreasing radius indicate the positions and approximate sizes of the progenitors of, respectively, Virgo, Fornax, the Milky Way and M31.

ature, the gas is fully ionized, suffers inefficient hydrogen cooling and can not form stars. This temperature criterion is widely used in hydrodynamical simulations (Stinson et al. 2006; Agertz et al. 2013). A common variation in simulations of the EoR, producing a similar effect, is to use a threshold in ionized fraction, as in Iliev et al. (2007, 2011); Ocvirk et al. (2014); Gillet et al. (2015).

The star particle mass at birth depends on the cell gas density, but is always a multiple of a fixed elementary mass

M_{\star}^{birth} , chosen to be a small fraction, $\sim 5\%$ of the baryonic mass resolution. In this framework, with the box size and resolution of CoDa (see Sec. 2.2), we have $M_{\star}^{\text{birth}} = 3549 \, M_{\odot}$. This mass is small enough to sample adequately the star formation history of even low mass galaxies, and still large enough to mitigate stochastic variations in the number of massive stars per star particle.

For each star particle, we assume that $\eta_{\text{SN}} = 10\%$ of its mass is in the form of massive stars which will go super-

nova after a lifetime $t_* = 10$ Myr, leaving no remnant. We consider a supernova energy $E_{SN} = 10^{51}$ erg per $10 M_\odot$ of progenitor, and this feedback was implemented in the RAMSES code using the kinetic feedback of [Dubois & Teyssier \(2008\)](#). After the massive stars have exploded, the remaining low mass stellar population is represented by a long-lived stellar particle of mass

$$M_* = (1 - \eta_{SN})M_*^{birth} = 3194M_\odot. \quad (2)$$

No chemical enrichment was implemented: indeed, extrapolating the results of [Maiolino et al. \(2008\)](#) to $z = 6-10$ yields very low metallicities for the gas ($Z/Z_\odot = 1/100 - 1/1000$), in a similar range to the simulations of [Pallottini et al. \(2014\)](#). At these metallicities, the cooling rates are still rather close to those for metal-free gas in most of the temperature range ([Sutherland & Dopita 1993](#)) and therefore we do not expect chemical enrichment to make a dramatic impact on our results.

In cosmological simulations, it is customary to rely on subgrid models, providing an effective EoS that captures the basic turbulent and thermal properties of the gas in turbulent, multiphase, centrifugally supported disks. Models with various degrees of complexity have been proposed in the literature, for instance using a polytropic EoS ([Yepes et al. 1997](#); [Springel & Hernquist 2003](#); [Schaye & Dalla Vecchia 2007](#)). We refrain from using this approach, because these subgrid models account for star formation and radiation from massive stars in star forming complexes, which are already explicitly accounted for in our feedback models. Therefore we set the RAMSES parameters so that the polytropic EoS does never kick in.

To summarize, we used for this simulation rather standard galaxy formation recipes, which have proven quite successful in reproducing various properties of galaxy evolution ([Ocvirk et al. 2008](#); [Governato et al. 2009, 2010](#)). The main novelty of this study is the inclusion of the coupling between the hydrodynamics and the radiation produced by the stars, treated by our radiative transfer module ATON, described in Sec. 2.1.2.

At the galactic scale, since galaxies can be smaller than one CoDa cell, their ISM is not resolved. However, the goal of our star formation model is to produce sources with suitable number density and ionizing photon output to reionize the box in a reasonable time, and we must calibrate it to do so.

The overall star formation efficiency and feedback effect that result from our modelling are subject to adjustment of the adopted efficiency parameters, δ_* and ϵ_* . While the previous experience of simulating galaxy formation with such prescriptions is a useful guide, there is some additional tuning required to simulate the EoR. For a self-consistent simulation of the EoR, these local parameters for the internal efficiency of individual galaxies should be chosen so as to achieve whatever global efficiency over time is required to make the simulated EoR satisfy the known observational constraints on the EoR. We discuss some of these global EoR constraints in more detail in Sec. 3.2. Unfortunately, it is not possible to perform multiple runs of the size and resolution of the CoDa simulation, each with different values for these efficiency parameters, since the computational cost is prohibitive. Each such simulation would have to run all the way to the end of reionization, and one cannot substitute for this by doing coarse-grained, lower-resolution versions of this

large volume to reduce the computational cost enough per simulation, since lower resolution simulations do not yield the same outcome as the CoDa simulation with its higher resolution. In practice, we solve this problem, instead, by performing a large suite of smaller-box simulations (e.g. as small as $4 h^{-1}\text{Mpc}$ on a side) with the same particle mass and grid resolution as CoDa but much smaller volumes, since each of these can be run on the same computer as CoDa but require many fewer CPU-GPU hours. In this way, we were able to adjust the efficiency parameters for CoDa in advance, to best guess the outcome of the much larger-volume simulation.

There are limitations that are inevitable in this procedure, however. For example, small boxes tend to reionize much more quickly (i.e. over a more narrow redshift range) than do large-boxes and to have fewer rare haloes, thereby delaying the first appearance of sources. Periodic boundary conditions for a small box also tend to make its reionization history differ from that in a large box, since, e.g., radiation from sources inside the box that escapes from it must re-enter that box, but there is no realistic accounting of more distant sources. In addition, cosmic variance makes it difficult to guarantee that the outcome of a random realization of Gaussian-random-noise initial conditions in such a small volume predicts the outcome of a different realization in a much larger volume. Finally, current observational constraints on the global EoR are still highly uncertain and continue to change with time as more data becomes available, and these do not, themselves, exclude the possibility that the reionization of the local universe differed in some way from that of the universe-at-large.

2.1.2 ATON

ATON is a UV continuum radiative transfer code, which relies on a moment-based description of the radiative transfer equations and uses the M1 closure relation ([González et al. 2007](#)). It tracks the out-of-equilibrium ionisations and cooling processes involving atomic hydrogen ([Aubert & Teyssier 2008](#)). Radiative quantities (energy density, flux and pressure) are described on a fixed grid and evolved according to an explicit scheme under the constraint of a Courant-Friedrich-Lewy condition (hereafter CFL), i.e. timesteps must be no larger than some fraction of a cell-crossing time for a signal travelling at the characteristic speed for radiative transport. In the case of interest here, involving ionization fronts in the low-density IGM, this speed can approach the speed of light. Each stellar particle is considered to radiate for one massive star lifetime $t_* = 10$ Myr, after which the massive stars die (triggering a supernova explosion) and the particle becomes UV-dark. We used a mono-frequency treatment of the radiation with an effective frequency of 29.61 eV for a 10^5 K black body spectrum as in [Aubert & Teyssier \(2008\)](#). This corresponds to a $\sim 100M_\odot$ pop III star, although the mass scale of such primordial stars is currently very much debated (cf the introduction of [Jeon et al. \(2014\)](#) for a summary). Finally, we assume each star particle to have an intrinsic emissivity of 4800 ionizing photons/Myr per stellar baryon (i.e. 1.824×10^{47} photons/s/ M_\odot), as in [Baek et al. \(2009\)](#), although the latter used a 50 000 K spectrum. Here the exact temperature of the stellar sources is actually of little importance: this temperature affects the effective ion-

izing photon frequency, which sets the effective cross-section of the photo-ionization reaction σ_E , and therefore the penetration of UV photons. In such EoR simulations, all UV, H-ionizing photons are consumed within a few cells of the ionization front (hereafter I-front), and variations of the UV photons' energy mostly affects the thickness of the I-front, but not its average position. Therefore, the emissivity of the stellar sources turns out to be more important than the effective frequency. Using a 50000K black body spectrum would yield results rather similar to the 10^5 K we used here, provided the emissivity in ionizing photon number is kept constant.

Because of the relatively high spatial resolution of the CoDa simulation for an EoR simulation, we do not make any correction in terms of a clumping factor, to account for small-scale inhomogeneity in the gas density that was unresolved by the grid spacing, as was done for the largest boxes of [Aubert & Teyssier \(2010\)](#). A price which must be paid for such a high spatial resolution, however, is that the CFL timestep limiter of the light-crossing time per cell is exceedingly small and would normally make the number of timesteps required to span the evolution of the system over a given physical time prohibitively large if all of the gas dynamical quantities had to be advanced with this same small timestep. The sound-crossing and gravity time-scales are orders of magnitude larger than this radiation transport time, however, so there is a great mismatch between the number of timesteps required to account accurately for mass motions and that required to evolve the radiation field and ionization state of the gas. We have solved this computational problem without sacrificing accuracy by separately advancing the radiative transfer and ionization balance rate equations on GPUs while simultaneously advancing those for the gas and gravitational dynamics on the CPUs. By programming those GPUs so that each performs hundreds of these smaller radiative transfer timesteps in the same wall-clock time as it takes the CPU to perform one dynamical step, we are able to overcome this computational barrier. ATON has been ported on multi-GPU architecture, where each GPU handles a Cartesian sub-domain and communications are dealt with using the MPI protocol ([Aubert & Teyssier 2010](#)). By achieving an x80 acceleration factor compared to CPUs, the CFL condition is satisfied at high resolution within short wall-clock computing times. As a consequence, no reduced speed of light approximation is required.

We then have to determine an escape fraction for our stellar particles, i.e. the fraction of UV radiation able to escape the stellar birth cloud represented by the particle, into the interstellar medium (hereafter ISM) of the cell it belongs to, which is different from the escape fraction from the entire galaxy. Here we discuss two potential sources of absorption in the ISM: Hydrogen, and dust. We will therefore consider the two corresponding escape fractions, $f_{\text{esc},*}^{\text{H}}$ and $f_{\text{esc},*}^{\text{dust}}$ (we use the $*$ subscript to remind the reader at all times that this escape fraction applies to a stellar particle and its cell and not to a galaxy and the IGM). Their product gives the stellar particle escape fraction $f_{\text{esc},*} = f_{\text{esc},*}^{\text{H}} \times f_{\text{esc},*}^{\text{dust}}$. While there is some guidance (although somehow controversial) in the literature for the choice of the escape fraction, we can consider two limiting cases:

- the large scale, low spatial resolution regime, where the

radiative transfer grid cells size is larger than the galaxies, e.g. $dx = 0.44 \text{ h}^{-1}\text{Mpc}$ as in [Iliev et al. \(2014\)](#). This regime prescribes the use of a “galactic escape fraction”, which accounts for the ionizing photons lost to the unresolved interstellar medium of each galaxy. In this case, $f_{\text{esc},*}$ can span a wide range of values, from almost 0 to about 40%, as shown by high resolution simulations ([Wise & Cen 2009](#); [Razoumov & Sommer-Larsen 2010](#); [Yajima et al. 2011](#); [Wise et al. 2014](#)).

- the very high resolution regime, which is closer to resolving the interstellar medium and star forming molecular clouds, such as [Kim et al. \(2013\)](#); [Kimm & Cen \(2014\)](#); [Rosdahl et al. \(2015\)](#), where the cell can be as small as $dx \sim 1 \text{ pc}$ in the most refined regions. In this case the totality of the stellar photons produced reaches the interstellar medium. Therefore the stellar particle escape fraction $f_{\text{esc},*} = 1$ although the *galactic* escape fraction can be much smaller.

In CoDa, as galaxies form, their progenitor dark matter haloes span volumes from several to hundreds of grid cells across. Even after they are fully collapsed, these haloes are still typically a few to a few tens of cells across. As a result, the effects of internal radiative transfer of UV photons through the interstellar hydrogen are accounted for, although in a rather coarse way. Hence, we set $f_{\text{esc},*}^{\text{H}} = 1$, so as not to over-count the attenuation caused by the interstellar hydrogen gas inside the galaxy. Dust may also play an important role: according to [Gnedin \(2014\)](#), dust optical depths as high as 0.8 in the UV could be common within galaxies by the end of the EoR, which amounts to a throughput of about $f_{\text{esc},*}^{\text{dust}} = 0.5$ of the UV flux. With these assumptions, we obtain a stellar particle escape fraction of $f_{\text{esc},*} = 0.5$.

Finally, we neglect any possible Active Galactic Nuclei (hereafter AGN) phase of our galaxies. Such sources could already be in place in rare massive proto-clusters during reionization ([Dubois et al. 2012b,a](#)). They are very rare and thought to be minor contributors to the cosmic budget of hydrogen-ionizing photons ([Haardt & Madau 2012](#); [Haardt & Salvaterra 2015](#)), although they could be important for explaining the line of sight variations of the properties of the Ly α forest just after reionization ([Chardin et al. 2015](#)).

2.1.3 Radiation-hydrodynamics coupling

We developed an interface that enables data exchange between RAMSES and ATON on the fly, leading effectively to a coupling between dynamics (handled by RAMSES) and radiation (handled by ATON).

At the end of a dynamical time-step, RAMSES sends to ATON the gas density, temperature and ionized fraction. Radiative transfer is then performed by ATON using these inputs, via sub-cycling of 100s to 1000s of radiative sub-steps. Once this sub-cycling is completed, ATON sends the temperature and ionized fraction back to RAMSES. However, CUDATON (the GPU port of ATON) can only handle regular Eulerian grids. For this reason, AMR is not used for this simulation, and RAMSES is effectively used as a hydro-PM code. Furthermore, CUDATON assumes a Cartesian domain decomposition for the MPI-parallelism when similar domains are treated by each GPU. Meanwhile, RAMSES relies on a space-filling curve decomposition. As a consequence,

data must also be reorganized at each transfer between the 2 applications.

This coupling method is similar to RAMSES-RT (Rosdahl et al. 2013) but with one photon group only and no AMR. RAMSES-RT has been extensively tested and passes the test suite of Iliev et al. (2009). RAMSES-CUDATON was similarly tested, as described in Stranex & Teyssier (2010)². As expected due to the similarity of the codes, RAMSES-CUDATON performs similarly well with these tests: various instances of Strömgren spheres and the photo-evaporation of a gas cloud, the hallmark of radiation-hydrodynamics coupling in galaxy formation.

One of the great advantages of RAMSES-CUDATON is its ability to run with the true speed of light, which stems from the combination of a moment-based method with GPU acceleration. This is quite unique among the handful of existing radiation-hydrodynamics codes. Some of them, mostly those involving ray-based photon propagation schemes, use infinite speed of light. Others, in particular those using moment-based methods such as ATON for the RT (e.g. RAMSES-RT), or Gnedin & Kaurov (2014), use the reduced speed of light approximation (as low as 1/100 of the real speed of light) in order to reduce the computational cost of the simulation. This is because in such a framework, the time-step of the code is set by the fastest physical process (gravitational, hydrodynamical and radiative). Since the speed of light is typically 100 to 1000 times faster than the speed of sound or bulk matter motions in cosmological simulations, the time-step must be about 100's of times shorter and therefore the simulation will be 100's of times slower than its pure hydrodynamical counterpart. While this is probably a good enough approximation in the dense interstellar medium of galaxies, it is not valid in the low density intergalactic medium, i.e. most of the volume of a cosmological simulation as we perform in the present paper (Rosdahl et al. 2013). Not only would the speed of ionization fronts in the IGM be misrepresented by the reduced speed of light approximation, but long-range feedback effects, like the impact of one massive object or a distant cluster of objects on the properties of low-mass galaxies may be impossible to capture in such a "slow light" framework. In RAMSES-CUDATON this problem is alleviated thanks to the GPU optimization. The x80 boost almost cancels out the added cost of the RT, and allows us to work efficiently with the real speed of light.

2.2 Simulation setup

2.2.1 Initial conditions

The initial conditions (hereafter ICs) were produced by the CLUES project (Constrained simulations of the local universe), assuming a WMAP5 cosmology (Hinshaw et al. 2009), i.e. $\Omega_m = 0.279$, $\Omega_b = 0.046$, $\Omega_\Lambda = 0.721$ and $h = 0.7$. A power spectrum with a normalization of $\sigma_8 = 0.817$ and $n = 0.96$ slope was used. The comoving box is 91 (i.e. $64h^{-1}$) Mpc on a side, with 4096^3 dark matter particles on 4096^3 cells and with the same cubic-lattice of 4096^3 cells for the

gas and radiation properties. The mass of each dark matter particle is then $3.49 \times 10^5 M_\odot$. The CoDa initial conditions have the average universal density for the chosen cosmology.

These ICs are tailored so as to reproduce the structure of the local universe at $z=0$, using constrained realizations of a Gaussian random field, in the spirit of Hoffman & Ribak (1991). In essence, the method takes as its input, galaxy observations of the local universe today: galaxy catalogs with distances and peculiar velocities, galaxy group catalogs, such as the MARKIII (Willick et al. 1997), SBF (Tonry et al. 2001) and the Karachentsev catalog (Karachentsev et al. 2004). These are used to infer the $z = 0$ dark matter distribution. Using the Hoffman-Ribak algorithm (Hoffman & Ribak 1991), a set of ICs (particle positions and velocities) at high redshift are produced, which will evolve into this present matter distribution over a Hubble time. This technique reproduces the large scale structure around the LG, although with some scatter, whereas small scales below $1 h^{-1} \text{Mpc}$ are essentially random. In order to make sure that a realistic LG is obtained, a second step was performed. A large number of similarly constrained realizations of the initial conditions, of order 200, was produced and a selection made for the one which best fit the LG and its relation to other known features of the local universe at $z = 0$, with masses and locations of objects like the LG and its largest galaxies, the MW and M31, and a cluster such as Virgo. The best realization is then enriched with higher order random k-modes to an arbitrary resolution, limited in practice by computer memory. The resulting ICs therefore contain:

- constrained low k-modes, driving the emergence of the large scale structures and galaxy clusters,
- unconstrained but "chosen" intermediate k-modes, picked from a sample of many realizations to produce a LG,
- fully random high k-modes, added by the resolution enhancement procedure.

This procedure is detailed in Gottloeber et al. (2010), and it produced the original CLUES ICs up to a resolution of 2048^3 in the full box. Here we have used Ginnungagap³ to increase further the resolution of the ICs up to that required by CoDa, i.e. 4096^3 .

A CoDa temperature map at $z=6.15$ in the supergalactic YZ plane is shown in Fig. 1. The white circles of decreasing radius denote the approximate locations and sizes of the progenitors of Virgo and Fornax galaxy clusters, M31 and the MW. Massive objects such as galaxy cluster progenitors appear as large bubbles of photo-heated gas in the temperature maps. Two galaxy group progenitors also appear on the map between the Local Group and Virgo, reminiscent of the CenA and M81 groups, as described for instance in Courtois et al. (2013). The exact occurrence, position and growth of these particular groups are not strongly constrained in our ICs generation procedure, but such groups are statistically expected to emerge in the filament that harbors the LG and that points to the Virgo cluster Libeskind et al. (2015). The impact of such groups on the formation and evolution of the LG could be quite significant and therefore their existence in the simulation is one of the virtues of using such tailored initial conditions.

² http://www.physikstudium.uzh.ch/fileadmin/physikstudium/Masterarbeiten/Stranex_2010.pdf

³ <https://github.com/ginnungagapgroup/ginnungagap>

Because of its location with respect to Fornax and Virgo, the LG shown in Fig. 1 can be considered as “the” LG of the box. While the fairly large size of the box is a requirement in order to reproduce well enough the local Universe without artifacts due to the periodic boundary conditions, and properly model reionization sources external to the Local Group, it has a third benefit: it will allow us find many galaxy pairs similar to the Local Group in mass ratios and separation through the box. This should therefore allow us to study “what if” scenarios, where analogs of the MW-M31 system would have evolved in a different environment.

The baryonic ICs were generated assuming a uniform temperature and the gas and dark matter are assumed to have identical velocity fields. The startup ionized fraction was set to the freeze-out value at $z = 150$. They are computed following standard recipes such as in RECFAST (Seager et al. 1999).

2.2.2 Code deployment

The CoDa simulation was performed on the Titan supercomputer at Oak Ridge National Laboratory. The computational domain spans a comoving cube 91 Mpc on a side, sampled on a fixed grid of 4096^3 cells. The code was deployed on 8192 cores, with each core coupled to 1 GPU, therefore requiring 8192 Titan nodes (1 GPU per node). Each node hosted one MPI process which managed a $128 \times 256 \times 256$ volume.

2.3 Run and data management

The run was performed from August to December 2013, and took about 10 days wall clock, using a total of about 2 million node hours, i.e. 60 million hours according to the Titan charging policy (1 node hour = 30 core hours). A total of 138 snapshots were written, every 10 Myr, from $z=150$ down to $z=4.23$, each one of size ~ 16 TBytes, summing up to more than 2 PBytes of data in total. Since it is not possible to save such a large amount of data in long-term storage, we decided to keep the following reduced data products:

- Full-box, low resolution (2048^3 grid) versions of the gas field (mass and volume-weighted temperature, volume-weighted pressure, density, velocities, ionized fraction), ionizing flux density, and of the dark matter density field, plus a list of all the star particles in the box, with their mass, coordinates and ages.

- Cutouts: for a set of 495 cubic subvolumes of $4 h^{-1}$ Mpc on a side, centered on regions of interest, the original, full-resolution data of the simulation, each with 256^3 cells, was saved. The coordinates of these cutout regions were pre-computed from a pure N-body simulation (no gas dynamics) by GADGET, with 2048^3 particles, using the same CLUES initial conditions, but coarsened for this lower resolution (i.e. instead of CoDa’s 4096^3 particles and cells), which we ran down to $z = 0$, to identify interesting regions in the present-day local universe. We will refer to this dark matter only simulation as DM2048. Each cutout contains all the gas quantities (same as full-box), the ionizing flux density, and the dark matter particles, at maximum resolution. The regions of interest we picked include the LG as determined by the CLUES constrained initial conditions. They also include 62 LG analogs, i.e. pairs of galaxies with masses, separations

and relative velocities compatible with the present-day MW-M31 system. They will be used in the future as examples of LGs forming in different environments. Each of these systems require 2 cutouts, one for the MW and one for M31. Finally, we added 399 regions following halo progenitors of various masses, from large galaxies to dwarfs, embedded in different environments, from cluster and galaxy groups to voids.

- Full box halo catalogs (see Sec. 2.4) containing the number of particles, position and velocity of each halo, as well as the list of dark matter particles it contains (ids and positions).

This strategy allowed us to bring down the data volume to a more manageable ~ 200 TBytes, while still allowing us to achieve our scientific goals.

2.4 Processing: friends-of-friends halo catalogs

We used the massively parallel friends-of-friends (hereafter FoF) halo finder of Roy et al. (2014) with a standard linking length of $b=0.2$ to detect dark matter haloes in the CoDa simulation. They are reliably detected down to $\sim 10^8 M_\odot$. More details on the resulting CoDa mass functions can be found in Appendix A.

2.5 Online data publication

We plan to make a subset of the data and higher level products publicly available through the cosmosim database hosted by Leibnitz Institut für Astrophysik Potsdam⁴ and the Vizier database at CDS Strasbourg⁵.

3 RESULTS

In this section we show that the simulation effectively captures the main physical processes we intended to describe, and show that the global behaviour of the simulation is correct. We emphasize the radiative-hydrodynamic processes, since it is the main novelty of this simulation, and finish by discussing the role of various halo mass ranges in the overall reionization process. Our intention is to introduce these results and highlight a few of them, while leaving further analysis and interpretation for future papers. Note: In what follows, in order to make comparisons possible between simulation results at a given redshift and some observable properties of the universe at a different redshift, we will, henceforth, refer to simulation redshift as z_{sim} , while z will refer to the observed, or “real world” redshift.

3.1 Maps: Zoom-In on Galaxies and the Cosmic Web during the EoR

3.1.1 Example: The M31 Progenitor Cut-out

The physical properties of the gas for the M31 progenitor cutout at 3 different epochs are shown in Fig. 2. The gas

⁴ <https://www.cosmosim.org/cms/simulations/cosmic-dawn/>

⁵ <http://vizier.u-strasbg.fr/viz-bin/VizieR?-source=VI/146>

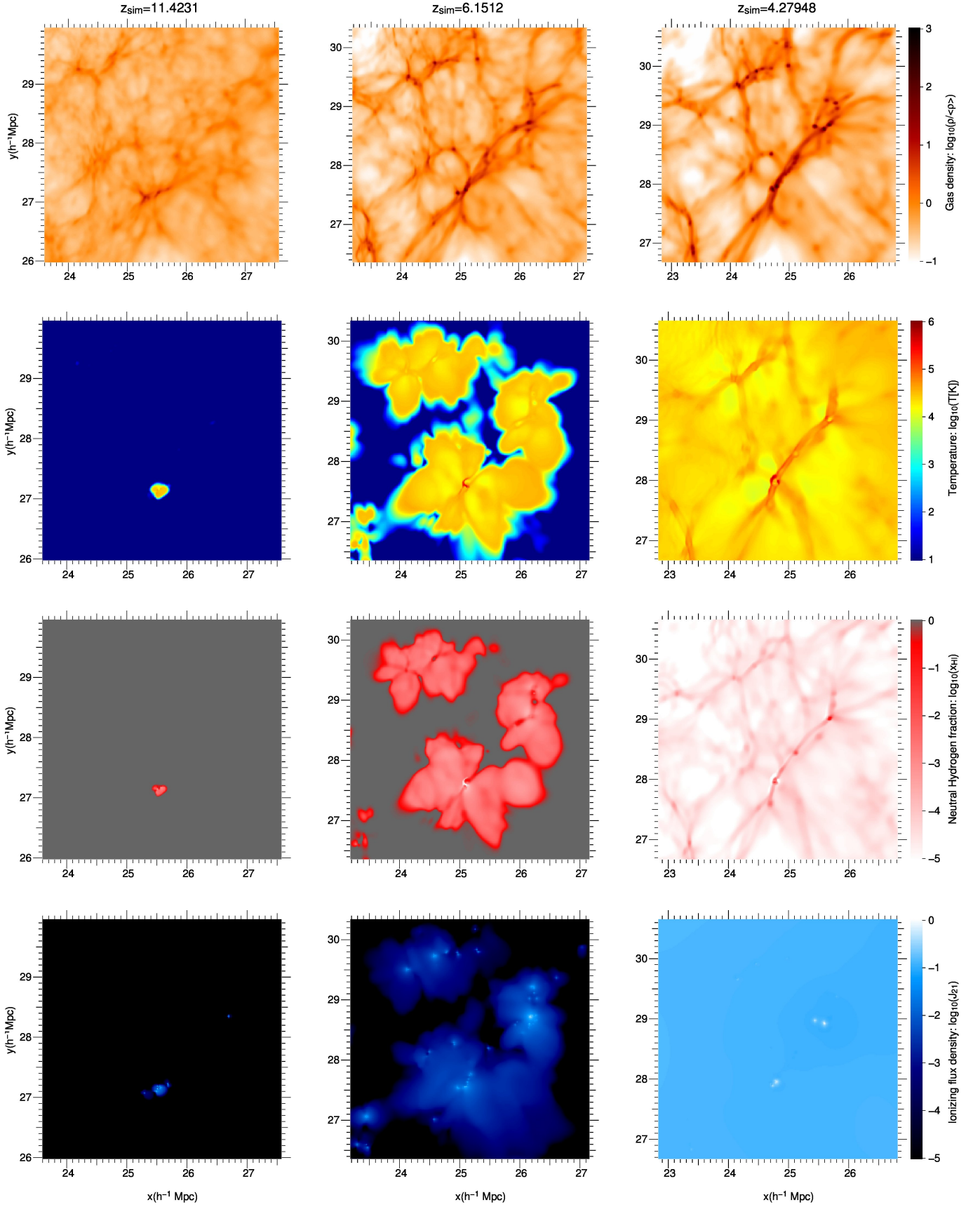


Figure 2. Maps of the physical properties of the gas in a slice through the comoving $4 h^{-1} \text{ Mpc}^3$ subvolume that contains the progenitor of M31. From left to right: $z_{\text{sim}} = 11$ to $z_{\text{sim}} = 4.28$ (end of the simulation). *Top:* co-moving gas density, *2nd row:* Temperature, *3rd row:* Neutral gas fraction, *bottom:* Ionizing flux density. Gas density and ionizing flux density are a projection based upon averaging over a slice 40 cells (or $625 h^{-1} \text{ kpc}$) thick, while the temperature and neutral fraction shown are taken from a slice just 1 cell thick (i.e. $15.6 h^{-1} \text{ kpc}$). (Madau & Furlanetto 2001)

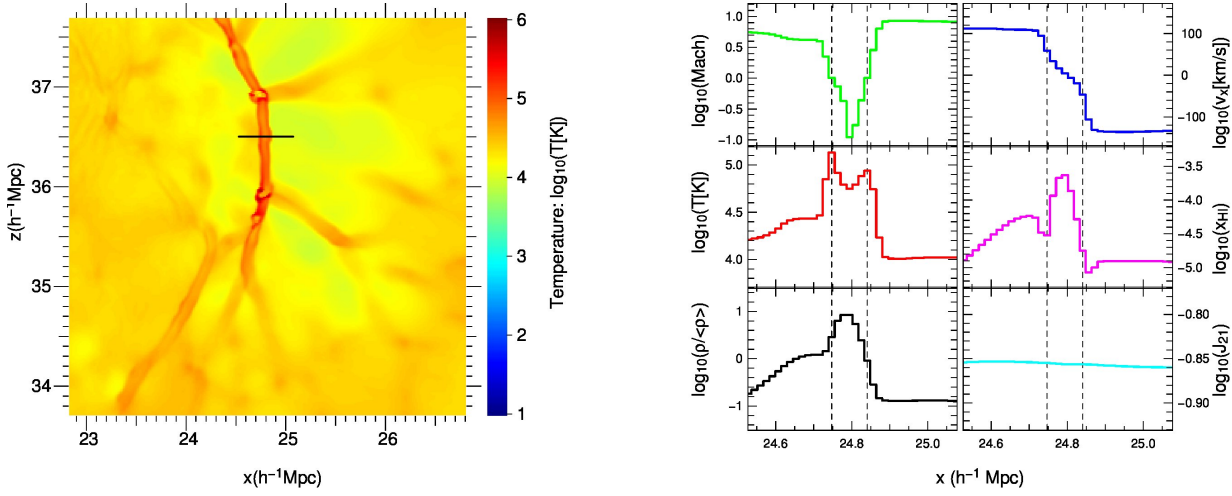


Figure 3. Illustration of the sheathed temperature structure of gas filaments. *Left:* a xz slab (1 cell thick) of the cutout of Fig. 2, chosen to highlight the filamentary structure. *Right:* gas overdensity, temperature, Mach number, x axis velocity, neutral Hydrogen fraction and ionizing flux density J_{21} along the black segment of the left panel. The vertical dashed lines mark the position of the two temperature peaks, for reference.

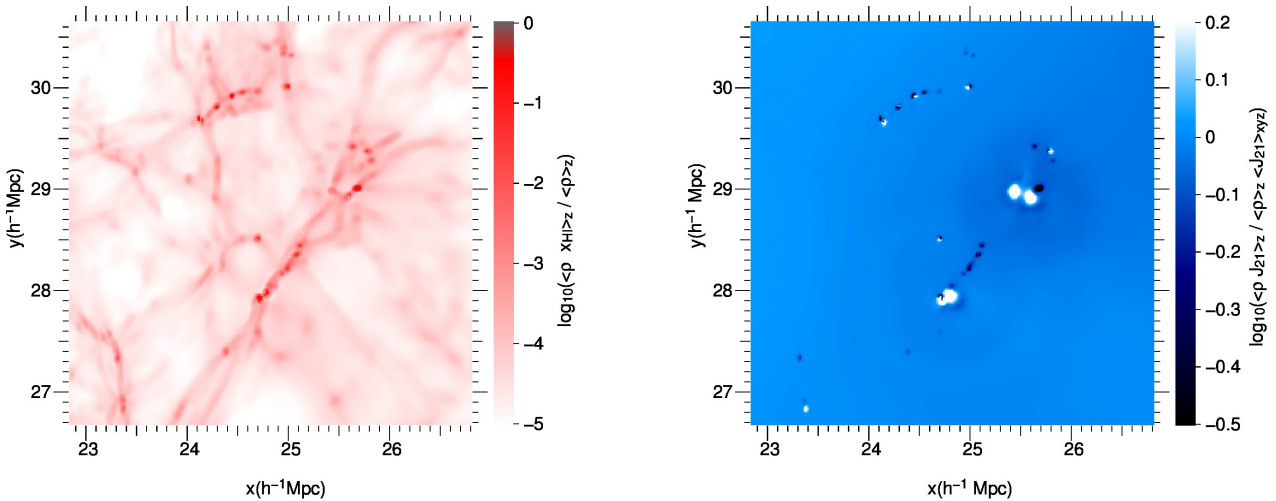


Figure 4. Mass-weighted maps (xy projection) of the neutral fraction (left) mass-weighted fractional excess of the ionizing flux density (right) for the M31 cutout at $z_{\text{sim}} = 4.28$. The fields are averaged over a 40 cells ($625 \text{ h}^{-1} \text{ kpc}$) thick slice.

density and ionizing flux density are averaged over a $625 \text{ h}^{-1} \text{ kpc}$ thick slice, while the temperature and neutral fraction shown are taken from a 1 cell thick slice (i.e. $15.625 \text{ h}^{-1} \text{ kpc}$). The gas is distributed along sheets, filaments and knots, which become more pronounced with time. The small scale structure visible in the high redshift maps progressively disappears, due to a combination of hierarchical merging and smoothing due to heating by UV radiation. Meanwhile, the voids get less dense, as their gas is pulled towards higher density regions. Let us recall that the filamentary structures seen in these maps are more often 2d planes extending perpendicular to the map than actual 1D filaments. Statistically, [Metuki et al. \(2015\)](#) showed that the volume in gas sheets is 7-8 times larger than the total volume of filaments, although the latter end up hosting a similar fraction of the total dark matter mass. Nevertheless we will call these filaments because that is what they look like in a 2D map, but

it should be kept in mind that these are often sheets rather than actual filamentary structures of gas.

In the high redshift cutout ($z_{\text{sim}} \sim 11$), the first star forms in the highest density clump, located at $(x, y) = (25.5, 27) \text{ h}^{-1} \text{ Mpc}$, and immediately starts to photo-heat its surroundings, as can be seen in the leftmost temperature map. Since the UV radiation is responsible both for ionizing the hydrogen gas and heating it to a few 10^4 K , it is not surprising that temperature, neutral fraction and ionizing flux density maps look very similar in structure. The ionization front propagates inside-out, faster in low density regions and slower in high density regions because the I-front speed is proportional to the number density ratio of ionizing photons to H atoms at the location of the front. This gives rise to the typical butterfly-shaped ionized regions seen in the middle temperature and neutral fraction maps. Radiation leaves the source's host halo, first into low-density regions,

and later on manages to penetrate denser nearby gas sheets and filaments.

The $z_{\text{sim}} = 6.15$ temperature map also demonstrates the very different scales of the feedback processes implemented: the yellow-orange regions are photo-heated, while the much hotter, reddest region near the highest density peak is also subject to supernova feedback. Therefore stellar formation in this simulation impacts the gas in two very different ways: UV radiation produces long range photo-heating, while supernovae form very local, but very hot bubbles. These supernovae-heated regions are significantly more highly ionized than regions which are only photo-heated.

The bottom middle panel shows that the photon propagation is restricted to photo-heated, ionized regions, as expected since the mean-free-path of ionizing UV radiation is very small in the neutral gas ahead of the I-fronts, so the boundaries of the H II regions are sharply defined. Additionally, this panel shows the location of young star particles, appearing as peaks in the ionizing flux density distribution. Their distribution follows closely the gas density peaks, in which they form. The last redshift (post-reionization) panel displays several interesting features, as follows.

3.1.2 Sheathed gas filaments

First of all, the temperature map shows that gas filaments after reionization (or after reionization is completed locally) are rather thick, as also seen for instance in Pawlik et al. (2009), although using a uniform background, and in Wise et al. (2014). Moreover, the filaments display a sheathed structure in temperature, with a hot, tube-like envelope surrounding a cooler core. This is also seen, for instance, in Ocvirk et al. (2008), although the aspect of the sheath looks thicker in the latter, perhaps because of the larger cell size in low density regions (i.e. the outskirts of the filament). More recently, Kaurov & Gnedin (2014) obtained sheathed filamentary structures very similar to CoDa's filaments in physical properties and thickness (about $\sim 200 \text{ h}^{-1} \text{ kpc}$ in both cases).

We investigated the origin of this structure. It could simply be due to the difference in cooling rates between the denser, more neutral core and the more diffuse filament envelope. There could also be a more dynamical process causing this temperature gradient: the filament could expand radially due to the outward pressure force caused by its photo-heating. During this expansion, its outer layers would get shocked, forming a high temperature layer similar to the contact discontinuity seen in the adiabatic shock tube experiment (Toro et al. 1994). To gain better insight into the process, we detail the properties of a gas filament in Fig. 3. The left panel shows an xz slab (one cell thick) of the cutout of Fig. 2, chosen to highlight the filamentary gas structure. The right panel shows the gas properties of individual cells along the short black segment crossing the filament of the left panel. The density is highest at the filament's center, and decreases outwards. The temperature shows a double peaked profile, surrounding the cooler, high density center. The temperature peaks, combined with the decreasing density, result in 2 dips in the Hydrogen neutral fraction. Moreover, the gas velocity field reveals that all of the gas is inflowing towards the filament, rather than out-flowing, and the temperature peaks are located at the position showing

the strongest velocity gradients, i.e. they demark a strong deceleration region in the gas, and the transition from a supersonic to subsonic flow, as shows the Mach number plot. Therefore, this hot layer is not a contact discontinuity (which would be flowing outwards) but rather an accretion shock formed by gas accreting onto the filament or plane. This is similar to the accretion shocks found for 1D pancake collapse by Shapiro & Struck-Marcell (1985), in which dark-matter-dominated gravity drives a cold, supersonic baryonic infall toward the central density peak, and this causes strong shocks that decelerate the baryons and convert their kinetic energy of infall to thermal energy (i.e. high T). It is also similar to the halo accretion shocks described by Birnboim & Dekel (2003). However, instead of spherical accretion as considered by Birnboim & Dekel (2003), we here have either planar, perpendicular accretion onto a sheet or radial, cylindrical accretion onto a filament.

3.1.3 Self-shielding

The filament's core is also significantly more neutral than its hot envelope. This is just the result of the higher density and therefore higher cooling and recombination rates: indeed, the bottom right panel of Fig. 3 shows that the ionizing flux density is completely flat across the filament: the UV background does not "see" the filament. It is not able to self-shield.

Here we investigate further the occurrence of self-shielding in CoDa. The low redshift ionizing flux density map of Fig. 2 shows that except at the location of the 3 most massive density peaks still forming stars, the pockets of radiation seem to have given way to an almost perfectly uniform UV background. To get a sharper view of this, we computed the ratio of the gas-mass-weighted ionizing flux density to the average J_{21} :

$$\langle \rho J_{21} \rangle_z / \langle \rho \rangle_z \langle J_{21} \rangle_{xyz}, \quad (3)$$

where the subscript z denotes an average in the z direction while $\langle J_{21} \rangle_{xyz}$ denotes the average value over the whole cutout volume. The resulting map is shown in the right panel of Fig. 4. Thanks to the mass-weighting, both *sources and sinks* appear. While the largest gas over-densities seem to host sources, the smaller ones tend to be dark and act as sinks. Interestingly, most haloes hosting sources seem to be also partly sinks, and the mass-weighted neutral fraction map shows that all photon sources are located next to rather neutral cells. This illustrates how galaxies are all at least marginally resolved in CoDa, i.e. even when they are just a few cells across, the code can still capture, albeit coarsely, their strongly heterogeneous internal structure.

A filament of gas joins the 2 main sources of the map. The lower half of it consists of a chain of small gas over-densities, and seems able to self-shield, as shows the photon-density map. On the other hand, the upper half is less dense, and does not appear self-shielded: its ionizing flux density is identical to the background. This link between density and self-shielding is made more obvious by the phase diagram (J_{21} versus gas density) shown in Fig. 5: the majority of the cells see a UV background at $J_{21} \sim -1$. The high density part of the diagram has 2 branches. In the upper branch, the cells are under the influence of a nearby source or contain one. The lower branch arises from self-shielding, and starts

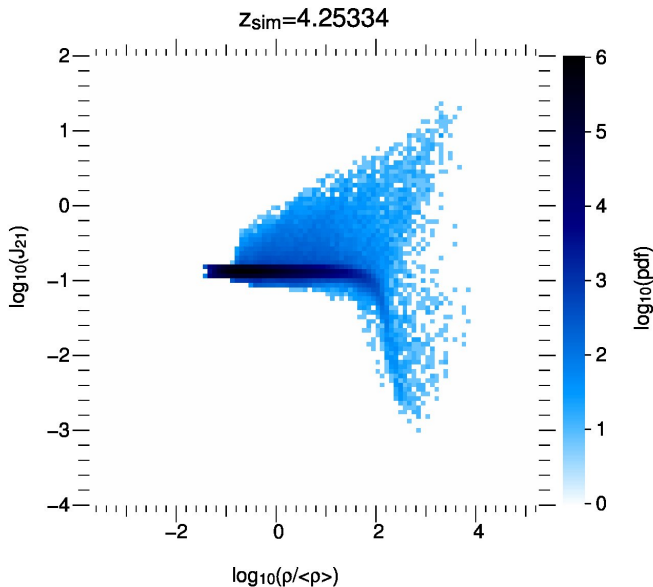


Figure 5. Distribution of the ionizing flux density versus gas overdensity for the M31 progenitor cutout, last snapshot.

at an overdensity of ~ 100 , similar to what was found by [Aubert & Teyssier \(2010\)](#). This value is at the very high end of the density range of filaments, and most filaments have lower densities, including that of fig. 3. Therefore, while haloes are at least partially self-shielded, filaments are not. Filaments can indeed resist the I-front propagation *at the onset of their own local reionization*, but not for long. The $z_{\text{sim}} = 6.15$ neutral fraction map indeed shows a few neutral clumps and filaments still partially neutral, surrounded by ionized gas and pockets of radiation. However this does not last and in the last snapshot they end up with neutral fractions between $x_{\text{HI}} = 10^{-2} - 10^{-4}$.

3.2 Global properties

In this section we show that the global evolution of the simulation is physically consistent and in line with our theoretical understanding of the EoR, and to some extent with the observational constraints available.

The most basic quantities to consider when gauging the success of a global EoR simulation are the evolution of the cosmic means of the neutral fraction, ionizing flux density, and cosmic star formation history. These are shown in Fig. 6, along with several observational constraints from [Fan et al. \(2006\)](#), [Ouchi et al. \(2010\)](#), [Calverley et al. \(2011\)](#), and the Planck CMB Thomson optical depth τ ([Planck Collaboration et al. 2016](#)). The observed cosmic star formation rate (hereafter SFR) constraints are taken from [Bouwens et al. \(2014\)](#): the grey area shows the envelope including the dust-corrected and dust-uncorrected SFRs. As a first remark, we note that the cosmic SFR in CoDa increases at all times, unlike the simulation of [So et al. \(2014\)](#), which shows a decline at late times. In the latter, the authors suggested this could be due to the small box size they used. This may indeed be the case, since CoDa is more than 95 times larger in volume.

The neutral fraction plot shows a characteristic very steep decrease, down to $x_{\text{HI}} \sim 10^{-4.2}$, where the slope be-

comes more gentle. This transition marks the end of the EoR, and correlates with the end of the surge in ionizing flux density seen in the middle panel. Following this definition, reionization is complete in the simulation at redshift $z_{\text{sim}} = 4.6$, which is about 300 Myr later than indicated by observational constraints. This discrepancy is consistent with the lower level of star formation in the simulation as compared to the observations of [Bouwens et al. \(2011\)](#), shown in the right panel.

These observational constraints are global ones, of course, and do not directly constrain the history of the local universe simulated by CoDa, to the extent that the local universe might be statistically differentiated from the global average.

In CoDa, the EoR lasts about 30% longer than global EoR observational constraints suggest. This could be a problem when using CoDa for producing mock observations of the Lyman α forest or the Gunn-Peterson trough for instance, for which the large size of the simulation is a desirable property. In an attempt to correct for this, we also show the effect of shrinking the redshift axis of the simulation: the solid black line in Fig. 6 shows the global properties of the CoDa simulation at the rescaled redshift $z_{\text{sim}} = z/1.3$. This simple rescaling mimics the effect of a modest increase of the star formation efficiency parameter. It brings the neutral fraction evolution in reasonable agreement with the observational constraints, as well as the ionizing flux density, the cosmic star formation history and the CMB Thomson scattering optical depth τ . Therefore the simulation should remain viable for producing mock observations, although other quantities, such as the collapsed mass fraction in halos, could be offset because of this temporal shift. We plan to explore this in future papers. In the rest of the paper, we carefully, explicitly state when such rescaled redshifts are used.

The post-reionization neutral hydrogen fraction in CoDa is about 0.5 dex lower than observed. Such an offset is not uncommon and the literature shows that simulations of the EoR can exhibit a variety of such small departures from the observed evolution of the ionized fraction measured from quasar lines of sight (see for instance [Aubert & Teyssier \(2010\)](#); [Zawada et al. \(2014\)](#); [Gnedin & Kaurov \(2014\)](#); [Aubert et al. \(2015\)](#)), in particular in the fully coupled radiation-hydrodynamics regime, due to the cost and difficulty of tuning such simulations. In the case of CoDa, this offset may be reduced by adopting a lower black-body temperature for our stellar sources, i.e. less massive stars, leading to a smaller effective ionizing photon energy, and therefore a smaller injection of energy per photo-ionization event, resulting in a lower average temperature of the large scale IGM, which would make it slightly more neutral. Indeed, Fig. 9 of [Pawlik et al. \(2015\)](#) shows that changing the heat injection per photo-ionization event has a strong impact on the post-reionization neutral fraction. However, decreasing the black body temperature of our sources to e.g. 50 000 K would also increase the effective H cross-section σ_E , leading to more ionizations and therefore may lead to an even larger ionized fraction at equilibrium. More simulations will be required to test this in future work. Finally, [Aubert & Teyssier \(2010\)](#) showed that modelling the gas clumping at sub-grid resolution could help in reproducing the observed trend. Just how much resolution is required to capture this additional gas clumping fully, beyond that of the CoDa sim-

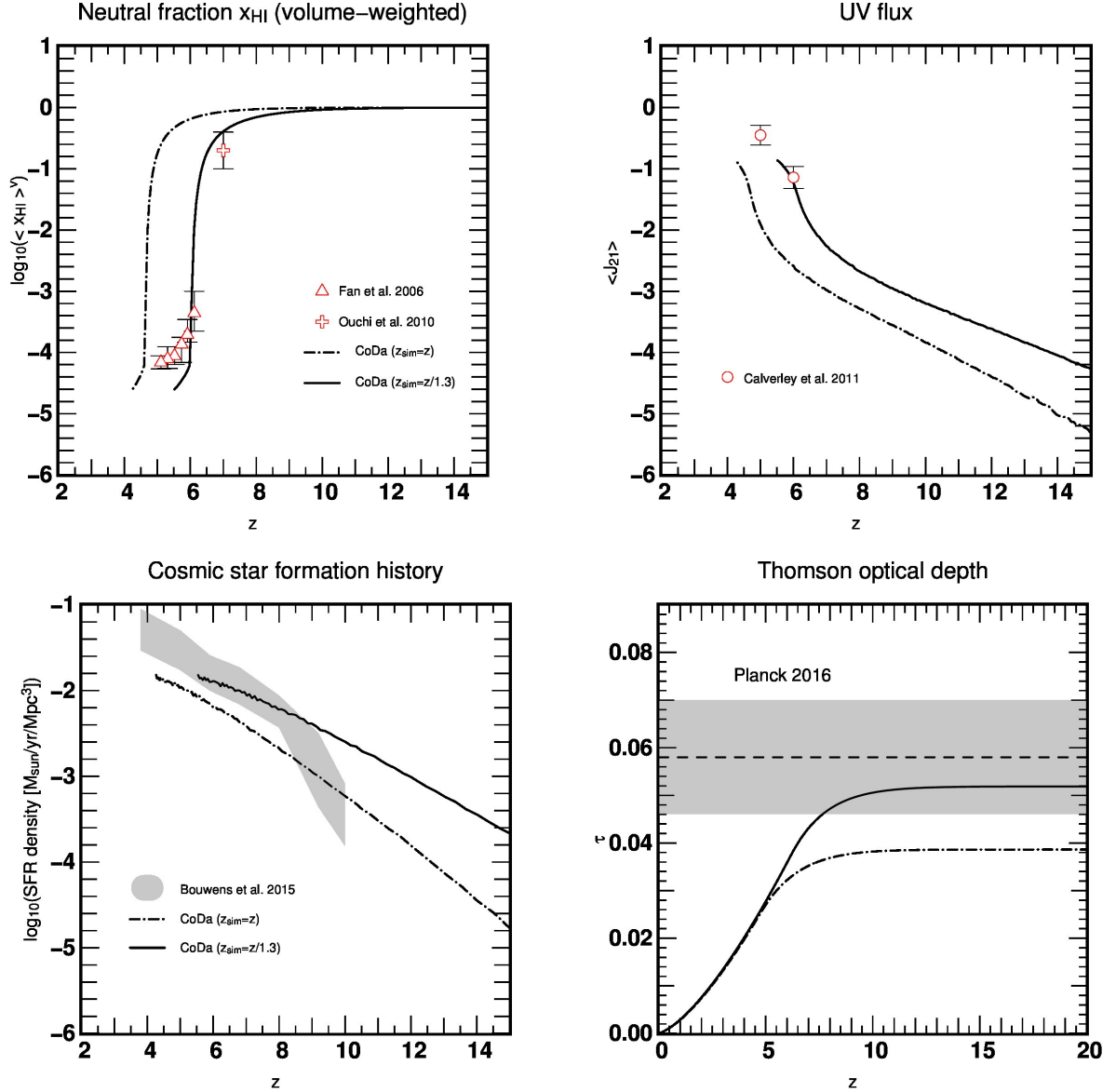


Figure 6. Global properties of the simulation, and comparison to observations. The dot-dashed black line shows the simulation results, while the solid black line shows a temporal rescaling of the CoDa simulation, i.e. for instance $x_{\text{HI}}(z_{\text{sim}} = z/1.3)$, so as to shift the end of the EoR to $z=6$ instead of $z=4.6$. This transformation brings all the fields in agreement with the available high- z observations.

ulation, was demonstrated, for example, by Emberson et al. (2013), who calculated the transfer of ionizing UV radiation through a static snapshot of the density field from a sub-Mpc-volume cosmological simulation. More recently, Park et al. (2016) used a fully-coupled radiation-hydro simulation of such a small volume to show that hydrodynamical back-reaction must be included, as well, since this causes the clumping to be time-dependent; eventually, even the denser, self-shielded gas ionizes and photoevaporates.

Large simulations such as CoDa, which can be run only once because of their cost, are very difficult to calibrate. As described in Section 2.1, for instance, tuning the subgrid star formation efficiency parameters by performing a large suite of small-box simulations (e.g. $4 h^{-1}\text{Mpc}$ boxes) to find parameters that satisfy observational constraints does not

guarantee that a large 91 Mpc box will follow the same evolution, because of cosmic variance, and because the large box will contain voids which can be orders of magnitude larger in volume than the small $4 h^{-1}\text{Mpc}$ calibration run. Moreover, it is unclear how periodicity (photons exit through one face and come back through the opposite face) affects the calibration of such small boxes. Taking a step back, it now appears that the strategy for calibrating such a large run should be more hierarchical. In the future, we can use the CoDa simulation, itself, to help re-calibrate, as well.

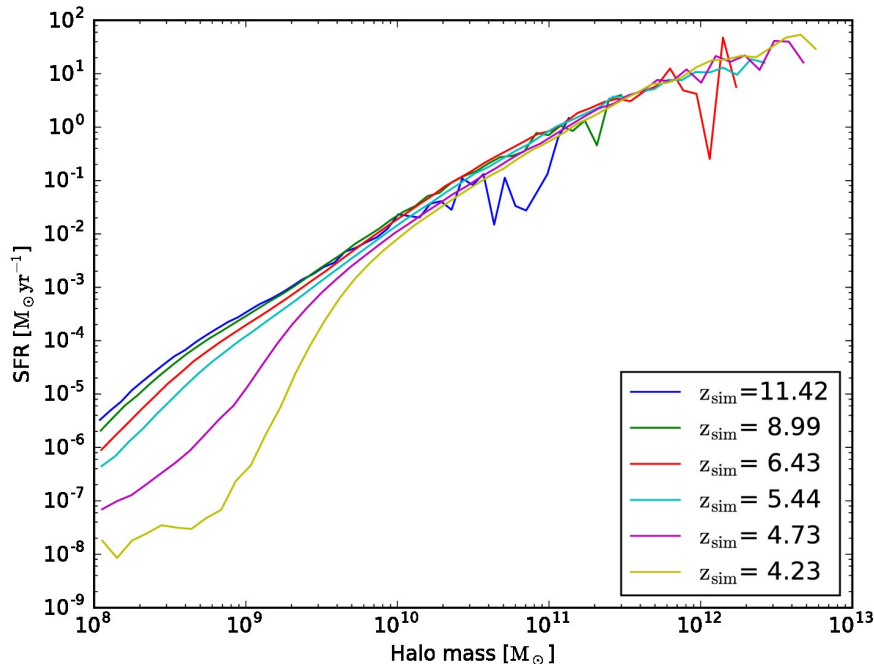


Figure 7. Instantaneous star formation rate per halo as a function of instantaneous halo mass, for various redshifts. The instantaneous SFR is computed as the stellar mass formed within an R_{200} radius sphere centered on the dark matter halo center of mass, during the last 10 Myr, divided by a duration of 10 Myr. Notice the sharp suppression at low mass.

3.3 Impact of radiative feedback on galaxy formation

In this subsection we investigate the impact of the radiation-hydrodynamics coupling on galaxy formation. Since early works of [Gnedin \(2000\)](#), several authors have addressed the impact of radiative feedback on galaxy formation, at low and high mass and in a number of contexts, with a fixed uniform UV background from [Haardt & Madau \(1996, 2012\)](#), and more recently with fully coupled RHD simulations ([Pawlik et al. 2013](#); [Wise et al. 2014](#); [Jeon et al. 2014](#); [Rosdahl et al. 2015](#); [Aubert et al. 2015](#); [Pawlik et al. 2015](#)), as a reduction and possibly a suppression of star formation at low masses. We computed the instantaneous SFR of CoDa haloes as the stellar mass formed within a sphere of radius, R_{200} (i.e. that within which the average density of the dark matter is 200 times the cosmic mean density) centered on the dark matter halo center of mass, during the last 10 Myr, divided by a duration of 10 Myr. Fig. 7 shows the instantaneous SFR that results, as a function of the instantaneous mass of the dark matter halo, for several redshifts.

There is a general trend for haloes at all masses to form fewer stars as time goes by, which is linked to a wide-spread decrease in accretion rate, as seen in [Ocvirk et al. \(2008\)](#). Here we find a trend: $\text{SFR} \propto M^\alpha$ for $M > 10^{10} M_\odot$ with a slope $\alpha \sim 5/3$. This slope is compatible with the values found in the literature. i.e $1 < \alpha < 2.5$, in numerical and semi-analytical studies, such as [Tescari et al. \(2009\)](#); [Hasegawa & Semelin \(2013\)](#); [Yang et al. \(2013\)](#); [Gong et al. \(2014\)](#); [Aubert et al. \(2015\)](#). However, the most striking feature of Fig. 7 is the very sharp decrease in SFR for the low

mass haloes, around $\sim 2 \times 10^9 M_\odot$. Before $z \sim 6$, low mass haloes sit on the same global trend as the high mass haloes. However, during the EoR, they transition from this “normal” state to a strongly suppressed state: at $z=4.2$, the $10^9 M_\odot$ haloes form almost 1000 times less stars than they did at $z=6$. This suppression reflects the great reduction of the gas fraction inside the galaxies that is below 20,000 K (as required for star formation by our star formation criterion described in Sec. 2.1.1), once the galaxy and its environment are exposed to photoionization during reionization. In contrast, the gas core of high mass haloes is dense enough to remain cool and/or cool down fast enough to keep forming stars, even if in bursts.

To check this, we ran two $8 h^{-1} \text{Mpc}$ RAMSES-CUDATON simulations: one with the exact same parameters and resolution (512^3 grid) as the CoDa simulation and a second without radiative transfer (therefore vanilla RAMSES, with SN feedback only), and no UV background. The star formation histories of haloes of various masses⁶ are shown in Fig. 8 for these 2 simulations. For the RHD simulation (left panel), there is again a very strong suppression of star formation at the end of the EoR for the low mass haloes. We note that more massive haloes are also affected, although in a less spectacular, if still significant way: even the $10^{10-11} M_\odot$ mass bin sees its star formation decrease by a factor of ~ 2 when compared to the pure hydro run (right panel). This experiment confirms that the addition of

⁶ the halo masses are those measured at the last snapshot, contrary to those of Fig. 7, which are instantaneous

radiative feedback via photoheating associated with hydrogen photoionization is the cause of the sharp suppression of star formation seen in low-mass haloes.

While this trend is clear and extends well up into the intermediate halo mass range, far above the minimum mass resolved in our halo mass function, the quantitative details may be affected by the limits of our spatial resolution. For instance, higher resolution simulations could form denser clumps inside a given halo, which would better resist ionization and photo-heating. Comparing with literature, we find for instance that O’Shea et al. (2015), using the renaissance simulations, offering up to 10 times higher mass resolution and better spatial resolution, but in a volume thousands of times smaller than CoDa, reports SFR suppression at halo masses typically 10 times smaller, i.e. $\sim 2 \times 10^8 M_\odot$. Meanwhile, Gnedin & Kaurov (2014), using the CROC simulations, find a very moderate impact of radiative feedback on the properties of low mass galaxies during reionization. However, they check for this effect by examining variations in the faint end of the UV luminosity function (hereafter LF). While this is valuable from an observational perspective, the impact of radiative feedback on star formation or the lack thereof is more easily assessed by looking at the SFR - halo mass relation as we do here, than at the UV LF. This difference in methodology makes a direct comparison with Gnedin & Kaurov (2014) difficult. Moreover, their treatment of star formation relies on modelling molecular hydrogen formation. This raises the density threshold for star formation, and therefore pushes star formation to higher mass haloes. The smallest galaxies, which are the ones most-strongly suppressed by radiative feedback in CoDa, would probably not have formed any stars at all within the CROC formalism. Therefore the question of the impact of radiative feedback on low mass haloes is tied, not only to spatial and mass resolution, but also to the choices made for the star formation recipe.

While performing a higher resolution CoDa would be prohibitively expensive with the current code and facilities, we performed a simple resolution study spanning 2 times higher and up to 4 times coarser spatial resolution with respect to CoDa, using 512^3 boxes of 4 to $32 h^{-1} \text{Mpc}$ on a side. Details can be found in Sec. B. Preliminary results indicate that degrading resolution beyond CoDa’s produces a radiative suppression of star formation at increasingly higher masses. However, the test boxes at CoDa resolution and 2 times higher resolution yield the same suppression mass scale of about $\sim 2 \times 10^9 M_\odot$. We note that this is compatible with the results of Pawlik et al. (2015), obtained with a mass resolution comparable to CoDa but better spatial resolution due to the use of an SPH-based Lagrangean hydro code instead of CoDa’s unigrid scheme. Their study reports a SFR suppression mass similar to ours, of $\sim 10^9 M_\odot$.

3.4 Contribution to the cosmic star formation density

We now investigate the contribution of haloes of various masses to the total star formation density of the box. While Fig. 7 already made clear that individual halo SFR increases with mass, one needs to multiply their SFR by the halo number density to obtain their contribution to the star formation density. This is shown in Fig. 9: each line shows

the star formation density due to each mass bin (each mass bin is a decade wide). As expected, the lowest mass haloes contribute almost equally to the other mass bins at high z ($z_{\text{sim}} > 11$) but their contribution decreases sharply during the late stages of the EoR, as they become increasingly suppressed. This decrease, though more gentle, is also seen for the $10^{9-10} M_\odot$ bin. However, for the 3 most massive bins, we find a constant rise of the star formation density. This rise is faster for increasing halo mass. This $10^{12-13} M_\odot$ mass bin can only be traced at $z_{\text{sim}} < 7$, as before this time the simulation does not contain any halo that massive. The $10^{10-11} M_\odot$ haloes dominate the cosmic star formation density for most of the simulation time, before being overtaken by the $10^{11} M_\odot$ mass bin. The hierarchy between mass bins and the overall evolution is very similar to that shown in Fig. 2 of Genel et al. (2014), although the details of the methodology and physics implemented differ vastly. The total SFR is however lower than the observational estimate, as was already pointed out in Sec. 3.2. This reflects the underestimated value adopted for the subgrid parameters which control the efficiency of star formation over-all, but not the relative contributions for different halo masses.

We will refrain, at this stage, from drawing conclusions as to which mass bin contributes most to large scale reionization. Indeed, while all stellar particles have the same intrinsic specific emissivity, we do not know how much of the ionizing photons produced made it into the IGM. We will come back to this question by measuring the circum-galactic escape fraction in a future study.

3.5 UV luminosity function

In order to gain more insight into the balance of bright vs faint galaxies in our simulation, we computed the luminosity function (hereafter LF) of CoDa haloes. We computed the M_{AB1600} magnitudes using the lowest metallicity stellar population models of Conroy & Gunn (2010). We ensured consistency with RAMSES-CUDATON’s source model by rescaling Conroy’s single stellar population models in flux so as to obtain the same ionizing photon output over 10 Myr. The results are shown in Fig. 10, along with observational constraints. The latter are taken from Bouwens et al. (2015) and Finkelstein et al. (2015), which have been shown to be in broad agreement with a number of other studies including Oesch et al. (2013) and Bowler et al. (2015). The LFs have been shifted vertically for clarity. The shaded area shows the envelope of the LFs obtained for 5 rectangular independent sub-volumes of the CoDa simulation. Each of these sub-volumes spans $\sim 150,000 \text{ Mpc}^3$ ($\Delta x = L_{\text{box}}/5$, $\Delta y = \Delta z = L_{\text{box}}$, i.e. $1/5$ of the full box volume), which is similar to the volume probed by CANDELS-DEEP at $z = 6$. The resulting envelope therefore illustrates the expected effect of cosmic variance at $M > -20$. The thick solid line shows the median of these 5 LFs. The observed $z = 6 - 8$ LFs are in rather good agreement with our simulation for $M > -21$.

However, the simulation seems to predict an overabundance of bright galaxies, at all redshifts. Due to the small simulated volume compared to the observations (the survey volume at the bright end is 4 times larger than CoDa), it is not clear if this is a robust prediction of the simulation or

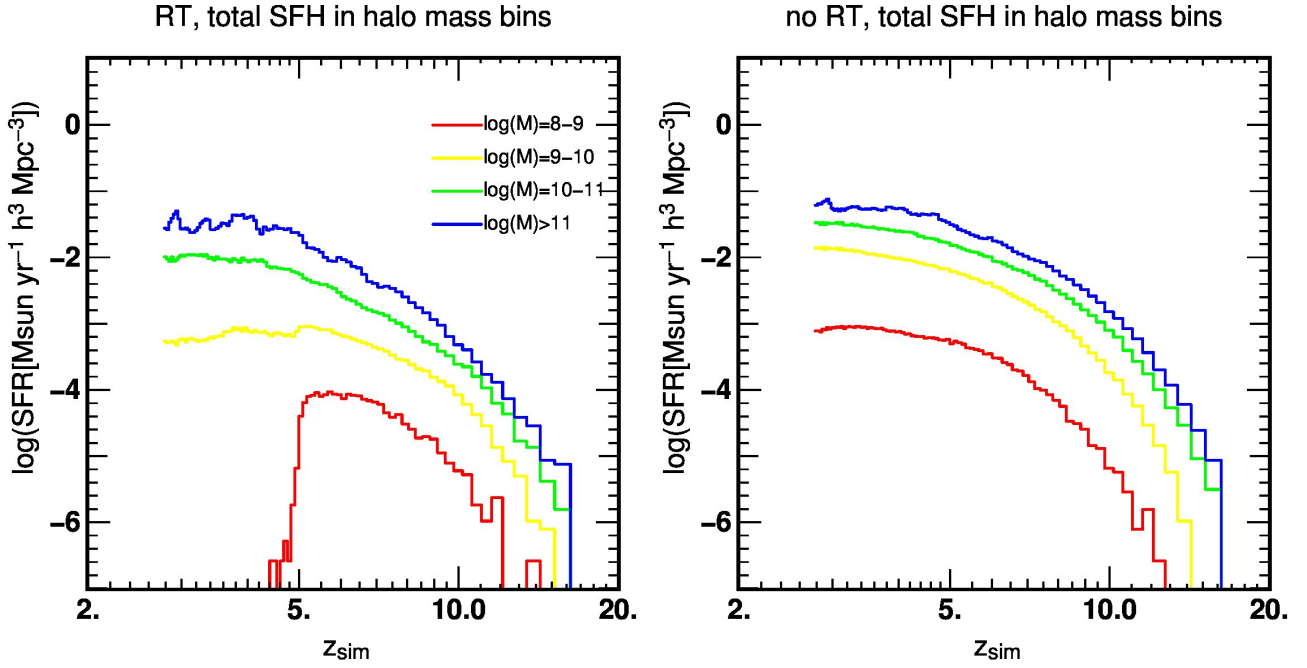


Figure 8. Total star formation histories of 4 halo mass bins for two simulations in a test box $8 h^{-1}$ Mpc on a side. *Left:* with full radiation hydrodynamics. *Right:* SN feedback only, no radiation. The mass bins correspond to the haloes' final mass (i.e. halo mass measured in the final simulation timestep, at $z_{\text{sim}} \sim 3$).

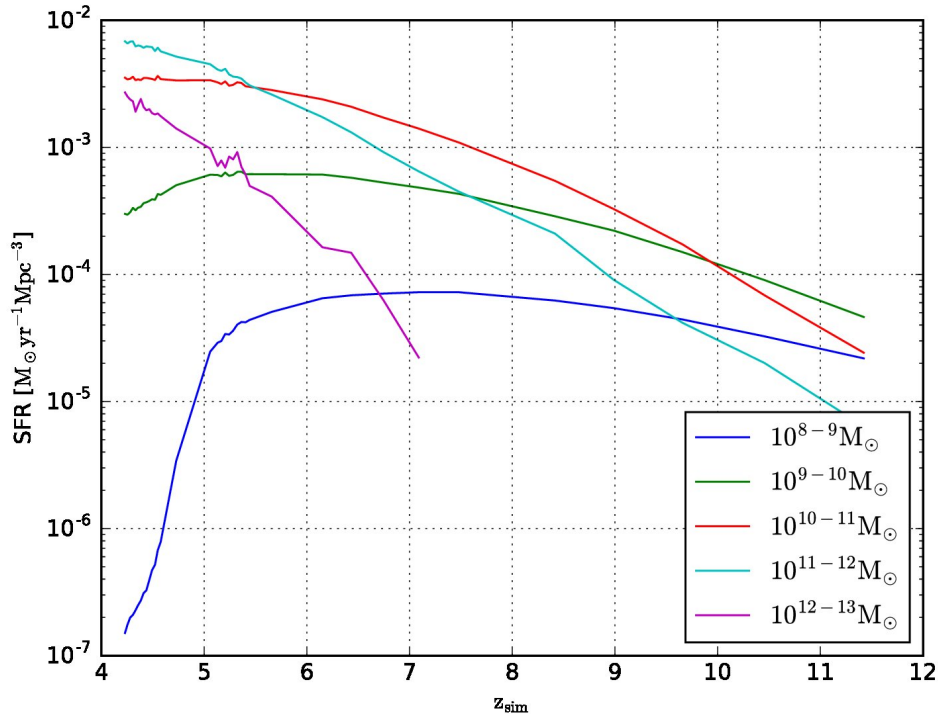


Figure 9. Cosmic SFR density as a function of redshift for various bins of instantaneous halo mass: each line shows the contribution of all galaxies within a given mass bin to the total cosmic SFR density.

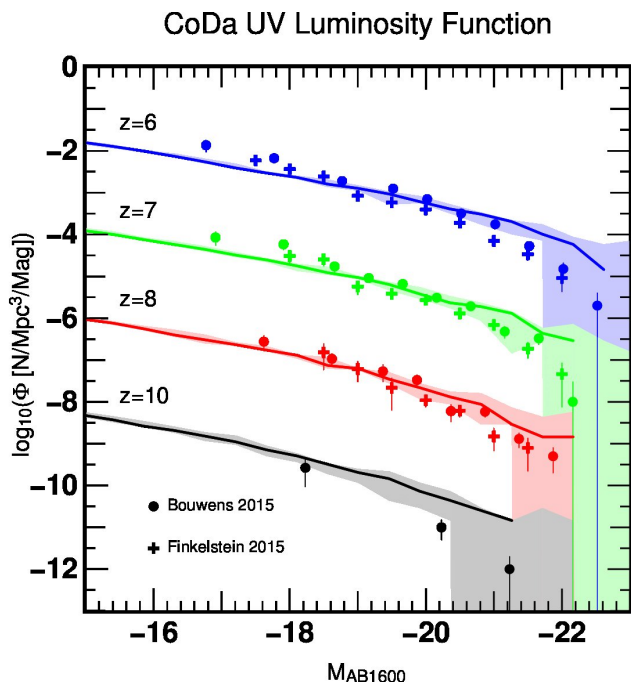


Figure 10. UV luminosity functions and comparison with observations. The full circles and crosses with error bars are the observations from Bouwens et al. 2015 and Finkelstein et al. 2015, respectively, at $z=[6, 7, 8, 10]$, while the shaded area and the thick line show the envelope and the median of the LFs of 5 equal, independent, rectangular sub-volumes taken in the CoDa simulation, at $z_{\text{coda}} = [6, 7, 8, 10]/1.3$. For clarity, the LFs have been shifted downwards by 0, 2, 4 and 6 dex.

just a statistical accident. There could however be several reasons for such an overabundance:

- Missing physics:

- dust: the only impact of dust extinction in CoDa is through a constant stellar escape fraction $f_{\text{esc},*}^{\text{dust}} = 0.5$. However, dust extinction becomes increasingly important at the bright end (Kimm & Cen 2013; Gnedin 2014). Moreover, observed high redshift galaxies, even during the EoR, could be fairly dusty (Watson et al. 2015).

- AGN feedback: the radiative and/or mechanical feedback from AGNs is believed to be responsible for the drop-off of the bright end of galaxy LF at low redshift. Could it be that early super-massive black holes in massive high redshift galaxies regulate the bright end of the galaxy LF as well? CoDa does not include AGN feedback, and it could help explain the overabundance in our LFs.

- overlinking: the FoF algorithm we used for halo detection is notorious for producing too massive haloes at high redshift, when compared to other halo-finding algorithms, such as the spherical overdensity. This is known as the “overlinking problem” (Watson et al. 2013). However, we checked that using a shorter FoF linking length $b=0.15$ instead of the usual $b=0.2$, which could be more adequate at high redshift, did not improve the LFs.

3.5.1 Halo Mass and UV luminosity:

We also investigated the relation between UV galaxy luminosity and halo mass. There is a general trend of increasing UV continuum luminosity with halo mass. This is shown in the left panel of Fig. 11, which corresponds to the galaxy population of the $z = 8$ UV LF of Fig. 10. Moreover, for a given halo mass, the luminosity can fluctuate significantly, as shows the vertical spread of the distribution. The vertical dispersion increases with decreasing halo mass, and is largest for the mass range sensitive to radiative feedback. The horizontal overdensity at ($M_{\text{AB1600}} \sim -10$, $M < 10^9 M_{\odot}$) is due to the quantization of stellar mass: the haloes located in this region have exactly one young UV-bright stellar particle, which has a fixed minimum mass. Below this limit, haloes are populated by star particles older than 10 Myr, therefore fainter in the UV. Indeed, in RAMSES-CUDATON, the mass of a stellar particle is always a multiple of the elementary star particle mass $M_{\star} = 3194 M_{\odot}$. This quantization of the stellar mass leads to a quantization in the magnitudes and therefore to possible artifacts in the faint-end LF, whereas the bright end LF is not affected, due to the larger stellar masses involved. We can mitigate the effect of this quantization by modifying the stellar particles’ masses by adding to them a random number taken from a uniform distribution between $[-M_{\star}/2, M_{\star}/2]$, where M_{\star} is the stellar particles elementary mass. A stellar particle of mass M_{\star} is therefore assigned a mass between $0.5M_{\star}$ and $1.5M_{\star}$, while a stellar particle of mass $10M_{\star}$ is assigned a mass between $9.5M_{\star}$ and $10.5M_{\star}$. Therefore this “smoothing” of the stellar particles’ mass is stronger, relatively, for haloes with lower stellar mass, but on average the total added stellar mass is 0. This simple procedure reduces quantization artifacts in the faint end of the LF, as shows the right panel of Fig. 11, and will help its interpretation.

4 DISCUSSION

4.1 Observing the End of Reionization: Depressing the Faint-End of the Luminosity Function

It has long been thought that reionization must have exerted a negative feedback on the ability of low-mass galactic haloes to form stars, by suppressing the infall of intergalactic baryons from the photoheated IGM following its local reionization (e.g. (Shapiro et al. 1994; Thoul & Weinberg 1996; Quinn et al. 1996; Kitayama & Ikeuchi 2000; Gnedin 2000)). As we mentioned in Section 1, this effect may help explain why there are so many fewer dwarf galaxy satellites observed in the Local Group than there are low-mass haloes predicted by N-body simulations of the CDM model. However, it has also been suggested that the suppression of star formation in low-mass galaxies by reionization feedback might be observable in the evolution of the global star formation rate and galactic luminosity function at high redshift. Barkana & Loeb (2000), for example, used a semi-analytical approach to show that if the local IGM Jeans mass filter scale jumped up when a patch became an H II region, a sharp drop in the global star formation rate would appear as reionization overtook a significant fraction of the volume of the universe. Large-scale N-body + radia-

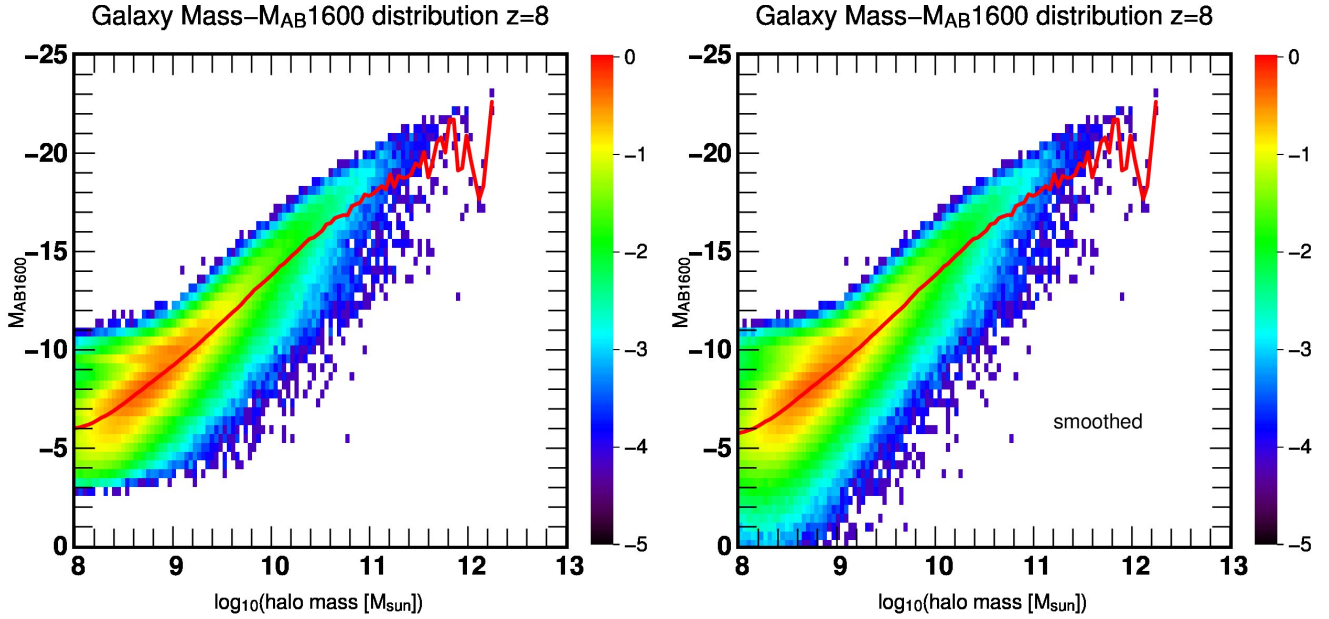


Figure 11. Galaxy mass - magnitude distribution at $z=8$ (i.e. $z_{\text{sim}}=6.15$). The color indicates the galaxy number density in $\text{N}/\text{Mpc}^3/\text{Mag}/\log(\text{M}_{\odot})$. The red line indicates the average Magnitude for each mass bin. Left: original stellar particle masses (i.e. quantized). Right: “smoothed” stellar particle masses.

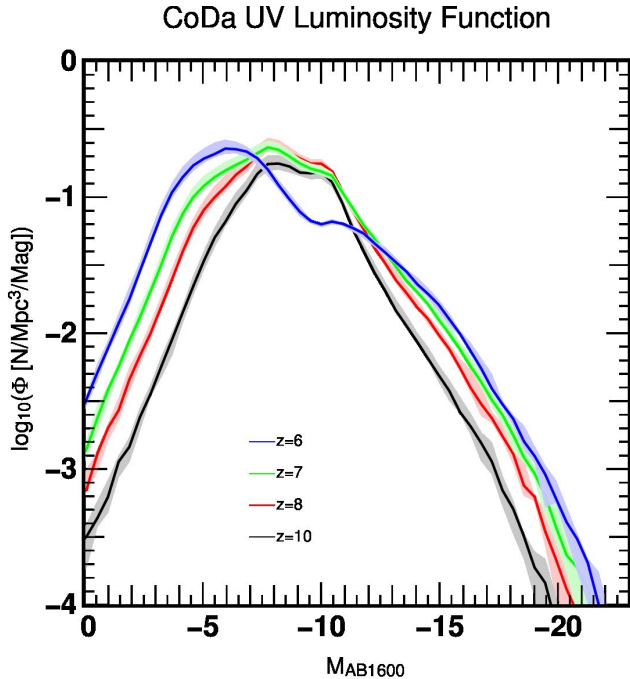


Figure 12. Same as Fig. 10, extended to include the faint end, using the “smoothed” stellar particle masses as explained in Sec. 3.5.1

tive transfer simulations of patchy reionization like those by Iliev et al. (2007, 2012), Ahn et al. (2012), and Dixon et al. (2016), for example, which also make this assumption about reionization suppression of low-mass galaxies inside H II regions find that reionization is, in fact, self-regulated,

such that the rise of the suppressible galaxies limits their ability to finish reionization. As a result, the dominant contribution, they find, comes from the larger-mass halos which are too big to be suppressed and whose number density is growing exponentially fast as reionization approaches completion. In that case, the global star formation rate does NOT decline sharply as reionization ends. However, the signature of this low-mass suppression may still be observable in ways that reflect the earlier onset and longer ramping up of reionization when low-mass galaxies dominate the early phase, before they saturate and the more massive galaxies dominate (e.g. Ahn et al. (2012); Park et al. (2013)). Since the low-mass suppressible dwarf galaxies whose contribution eventually drops off are also the ones at the faint end of the luminosity function, these expectations are consistent with another, possibly observable effect suggested by Barkana & Loeb (2000), of a drop in the galaxy LF at the faint end, marking the end of reionization. Similar conclusions have also been drawn more recently by, for example, Duffy et al. (2014).

We can use our CoDa simulation results to address both of these questions: Does the global star formation history show a strong drop as reionization approaches the end? And is there a depression of the faint end of the LF at this time?

According to our results for the global SFR density plotted in Fig. 6, there is no sharp downturn in the SFR as reionization runs to completion, which is not surprising, given the relative contribution to this SFR from haloes of different mass plotted in Fig. 9. The galaxies above $10^{10} \text{ M}_{\odot}$, which Fig. 7 indicates are basically too massive to be strongly suppressed by external radiative feedback alone during reionization, are found to dominate that SFR. Of course, we cannot say what the SFR history would have been if we had neglected radiative feedback. Our comparison in Fig. 8 of the SFRs in the two test box simulations (SN-feedback only, no

radiation, for one, and SN-feedback and radiative transfer in the other) as a function of the final masses of the haloes within which the stars reside at the end of the simulations at $z = 3$, shows that there can, in fact, be a depression of the SFR even in the higher mass bins when radiation is included, but much less dramatic than for the lower-mass haloes. And whether this is internal feedback or external reionization feedback is not evident from that plot alone. Nevertheless, it is a fact in the CoDa simulations that the global SFR continues to rise even up to the end of reionization, as seen in Fig. 6, and contrary to the expectations of Barkana and Loeb (2000). However, since radiative star formation suppression is very efficient in CoDa, it results in a change of shape of the faint end of the galaxy UV LF, as shown in Fig. 12, which uses the smoothed stellar particle masses, as explained in Sec. 3.5.1. The $M_{AB1600} = -10$ to -12 galaxies are hosted by haloes affected by radiative suppression by the UV background. As their SFR drops post-reionization, so does the abundance of galaxies in this magnitude range. Therefore, even though we find no sudden drop in the cosmic SFR density, we do see a sharp change in the shape of the LF at the very faint end, due to reionization, as proposed by e.g. Barkana & Loeb (2000). With respect to a similar exploration by Gnedin & Kaurov (2014), we find, like them, that the total cosmic SFR evolution is smooth throughout reionization because it is dominated by the more massive, radiation-immune galaxies, but unlike them, we find that reionization affects the shape of the faint end LF.

The bright end of the LF increases over time continuously, reflecting the convolution of the rising halo mass function with the halo-mass dependence of the star formation rate for haloes too massive to be suppressed by reionization. At fainter magnitudes but still brighter than -10 , the feedback from reionization on lower-mass haloes is reflected in the depression of the LF toward the end of reionization, in the magnitude range -10 to -12 . For magnitudes even fainter than -10 in Fig. 12, the LF reflects the passively evolving stars inside the low-mass suppressible haloes which have stopped forming new stars after they are overtaken by reionization. The section fainter than -10 shifts the peak to fainter magnitudes over time, as the stars formed inside low-mass haloes before they were fully suppressed get older.

4.2 The dark, not missing, satellites

The suppression of SFR we measured in low mass haloes is an important process in reducing the number of bright haloes, and therefore towards a remedy to the missing satellites problem. Indeed, Fig. 13 shows that the fraction of bright haloes (i.e. haloes hosting at least one stellar particle) is a steep function of the dark matter halo mass. Below $10^8 M_\odot$, more than 99% of the haloes are dark. The transition between dark and luminous haloes takes place between 10^8 and $10^9 M_\odot$. The transition mass (i.e. the mass for which half of the haloes are dark) does not evolve much prior to reionization but shifts quickly to higher masses near the end of the EoR and just afterwards ($z = 6$ and $z = 5.5$). This evolution is similar to that seen in Sawala et al. (2014), although the latter used an instantaneous and uniform reionization model, implemented as a uniform heating of the gas (i.e. no coupled radiation-hydrodynamics was performed as opposed to CoDa). Thanks to this simplification in the treat-

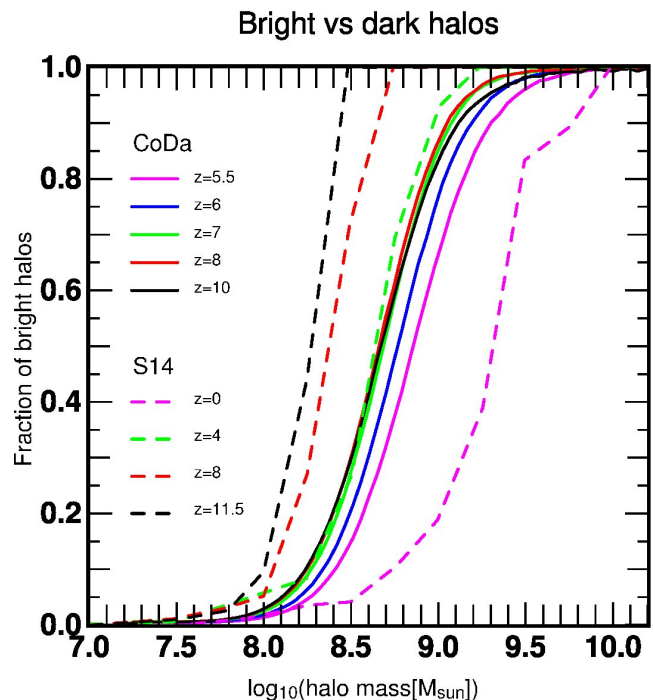


Figure 13. Fraction of bright halos as a function of mass and redshift. The original CoDa redshifts corresponding to the labels are $z_{\text{sim}} = [5.5, 6, 7, 8, 10]/1.3$, i.e. as in Fig. 10. The S14 curves show the results of Sawala et al. 2014.

ment of the EoR, they were able to carry out their simulation down to $z=0$ and show that such a mass-dependent reduction in the fraction of luminous haloes is able to quantitatively match the observed abundance of satellites (Sawala et al. 2015). Therefore, although CoDa did not run down to $z=0$, the trends we measured in the suppression of star formation in low mass haloes suggest that CoDa’s $z=0$ low mass halo population will be predominantly dark, alleviating the missing satellites problem.

5 CONCLUSIONS

CoDa (Cosmic Dawn) is a very large, fully coupled radiation-hydrodynamics simulation of galaxy formation in the local universe during the Epoch of Reionization. It was performed on Titan at Oak Ridge National Laboratory using RAMSES-CUDATON deployed on 8192 nodes, using 1 core and 1 GPU per node. This is the first time a GPU-accelerated, fully coupled RHD code has been used on such a scale. The simulation accurately describes the properties of the gas and its interplay with ionizing radiation, in particular the growth of typical butterfly-shaped ionized regions around the first stars and first galaxies, accompanied by photo-heating and the subsequent progressive smoothing of small scale gas structures. Gas filaments tend not to be self-shielded once the reionization radiation sweeps across them; the flux density of the ionizing radiation internal to these filaments thereafter is identical to that of the background. However, they are indeed slightly more neutral than surrounding voids. Furthermore, they develop a sheathed temperature structure, showing up as a hot tubular envelope

surrounding a cooler core, similar in nature to the more spherical accretion shocks seen around forming galaxies. On the other hand, haloes hosting gas denser than $100 \langle \rho \rangle$, when they do not host an ionizing source, show up as photon sinks in ionizing flux density maps, and are therefore self-shielded.

The low star formation efficiency assumed in our simulation leads to late reionisation. This can be corrected by a simple contraction of the redshift axis, designed to make reionization complete at $z = 6$. This remapping of the redshifts mimics the effect of a modest increase of the star formation efficiency parameter, and brings the simulation in agreement with several other observational constraints of the EoR, including the Thomson scattering optical depth measured by the Planck mission, the evolution of the ionizing flux density and the cosmic star formation rate.

The star formation rate of individual galaxies of a given mass is on average higher at high redshift and decreases as the Universe expands. Galaxies below $\sim 2 \cdot 10^9 M_\odot$ are strongly affected by the spreading, rising UV background: their star formation rate drops by a factor as high as 1000. This suppression reflects the great reduction of the gas fraction inside the galaxies that is below 20,000 K, once the galaxy and its environment are exposed to photoionization during reionization.

This produces a large number of sterile, low-mass haloes by the end of reionization, and a corresponding depression of the faint end of the UV luminosity function in the magnitude range $M_{AB1600} = [-12, -10]$. The latter is similar to an effect suggested by [Barkana & Loeb \(2000\)](#), although the accompanying sharp drop in the overall star formation rate they also suggested is not supported by the CoDa simulation results. This is explained by the fact that the total star formation rate is dominated at that time by more massive galaxies, too massive to be subject to this strong reionization feedback suppression.

The halo mass scale below which the CoDa simulation finds galactic halo star formation to be reduced by radiative feedback during the EOR was checked for numerical resolution effects by a set of smaller-box simulations of different mass resolution, described in Appendix B. The results showed that, while lower resolution than that of the CoDa simulation would increase the value of the threshold halo mass for suppression, a finer resolution by a factor of two in spatial resolution and eight in mass resolution (replacing one cell by 8 cells) finds the same threshold mass, suggesting that the value of the suppression mass is not an artifact of limited numerical resolution.

Although CoDa did not run down to $z=0$, the trends we measured in the suppression of star formation suggest that CoDa's $z=0$ low mass halo population would be predominantly dark, alleviating the missing satellites problem. In contrast, the gas core of high mass haloes is dense enough to remain cool and/or cool down fast enough to keep forming stars, even if in bursts. Overall, star formation in the whole box is dominated by $10^{10} M_\odot$ haloes for most of the EoR, except at the very end where $10^{11} M_\odot$ haloes become frequent enough to take the lead. The CoDa luminosity functions are in broad agreement with high redshift observations, except for the most luminous objects, where the number counts are subject to strong cosmic variance and may be affected by additional processes, such as evolving dust content and AGN feedback.

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APPENDIX A: CODA HALO MASS FUNCTIONS

The mass functions (hereafter MF) obtained with FoF (Roy et al. 2014) at $z_{\text{sim}}=10.28$ and $z_{\text{sim}}=4.41$ are shown in Fig. A1, along with a Sheth-Tormen MF (Sheth & Tormen 2002) and a FoF universal fit from Watson et al. (2013). At $z_{\text{sim}}=4.41$, the CoDa MF is fairly well represented by both fits for haloes larger than 1000 particles, but sits slightly above the Poissonian error bars of the Watson FoF MF (gray area) at the high mass end. To check the origin of this excess, we also plot the MF of the DM2048 simulation, a dark matter only companion simulation run with the N-body code Gadget 2 using the same CLUES initial conditions as CoDa degraded to 2048^3 resolution. This comparison is useful because DM2048 has about ~ 20 times better force resolution than CoDa: tree codes such as Gadget 2 (Springel 2005) perform well with a force resolution set to $1/20$ to $1/40$ of their average interparticle distance, while CoDa’s force resolution is equal to ~ 1.5 times the cell size (equal to the average interparticle distance), due to the unigrid scheme, as shown for instance in Fig. 1 of Teyssier (2002). The DM2048 can therefore be used as reference: it can inform us on possible artefacts due to CoDa’s unigrid gravity solving. The DM2048 MF also displays some excess at the high mass end, comparable to CoDa’s. This is confirmed by the bottom panel, showing the ratio of CoDa MF to DM2048 MF, which is close to 1 at the high mass end. In contrast, the MFs differ significantly at low masses, with CoDa showing an increasing deficit of low mass haloes compared to DM2048. We attribute this deficit to CoDa’s limited force resolution, which

| Box size ($h^{-1}\text{Mpc}$) | Grid size | Suppression mass ($10^9 M_{\odot}$) |
|------------------------------------|-----------|--|
| 4 | 512^3 | 1.7 |
| 8 | 512^3 | 1.7 |
| 16 | 512^3 | 8 |
| 32 | 512^3 | 40 |

Table B1. Parameters of the simulations of the resolution study. The gray row corresponds to the CoDa resolution. The suppression mass is defined as the intersection between the high mass SFR fit and the low mass SFR fit of each simulation.

hinders the proper resolution and detection of the smallest haloes. This deficit is present in the $z_{\text{sim}}=10.28$ MFs as well (left panels), with a similar amplitude. Above 1000 particles, on the other hand, CoDa displays an excess of haloes compared to DM2048. This is also likely caused by the limited force resolution of CoDa: the unresolved low mass haloes provide a large pool of “free”, untagged particles which the FoF may spuriously link to massive haloes and therefore increase their particle numbers. All in all, in order to mitigate these effects in our analysis, we will refrain from analysing haloes less massive than $10^8 M_{\odot}$. Above this mass, our MF is uncertain by no more than a factor of two on average, but much better in general, in particular above 1000 particles and at lower redshifts.

APPENDIX B: NUMERICAL RESOLUTION AND SUPPRESSION MASS

As described above, CoDa finds that star formation is suppressed in low-mass haloes by the feedback associated with reionization, for haloes in the range roughly below $\sim 2 \times 10^9 M_{\odot}$. To investigate the dependence of this suppression mass on the size and mass resolution of the simulation, we performed a series of smaller-box simulations, with input parameters identical to those of CoDa, but a range of resolutions, both higher and lower than that of CoDa. The simulation parameters are summarized in Table B, along with the resulting suppression mass obtained in each case. As seen in Table B, a fixed number of particles and cells are adopted for a hierarchy of different box sizes, so their (space, mass)-resolutions range from (2,8) to (1/4,1/64) times those of CoDa.

Since we kept the input parameters for star formation efficiency the same in all simulations, the higher resolution cases produced a higher star formation rate (by resolving higher density gas and thereby triggering the subgrid star formation criterion more often) and, hence, ended reionization earlier, as shown in Fig. B1. In order to compare the halo mass scales of suppression at different resolutions, therefore, it is necessary to make an adjustment for this displacement of the reionization time-histories relative to each other for the different cases. Fortunately, this is a well-defined operation. The global reionization histories in Fig. B1 all have in common a very sharp drop in the neutral fraction at the end of reionization, followed by a much flatter, slow decline thereafter in the post-reionization era, controlled then by the average UV background and IGM density. A well-defined epoch of comparison is that which corre-

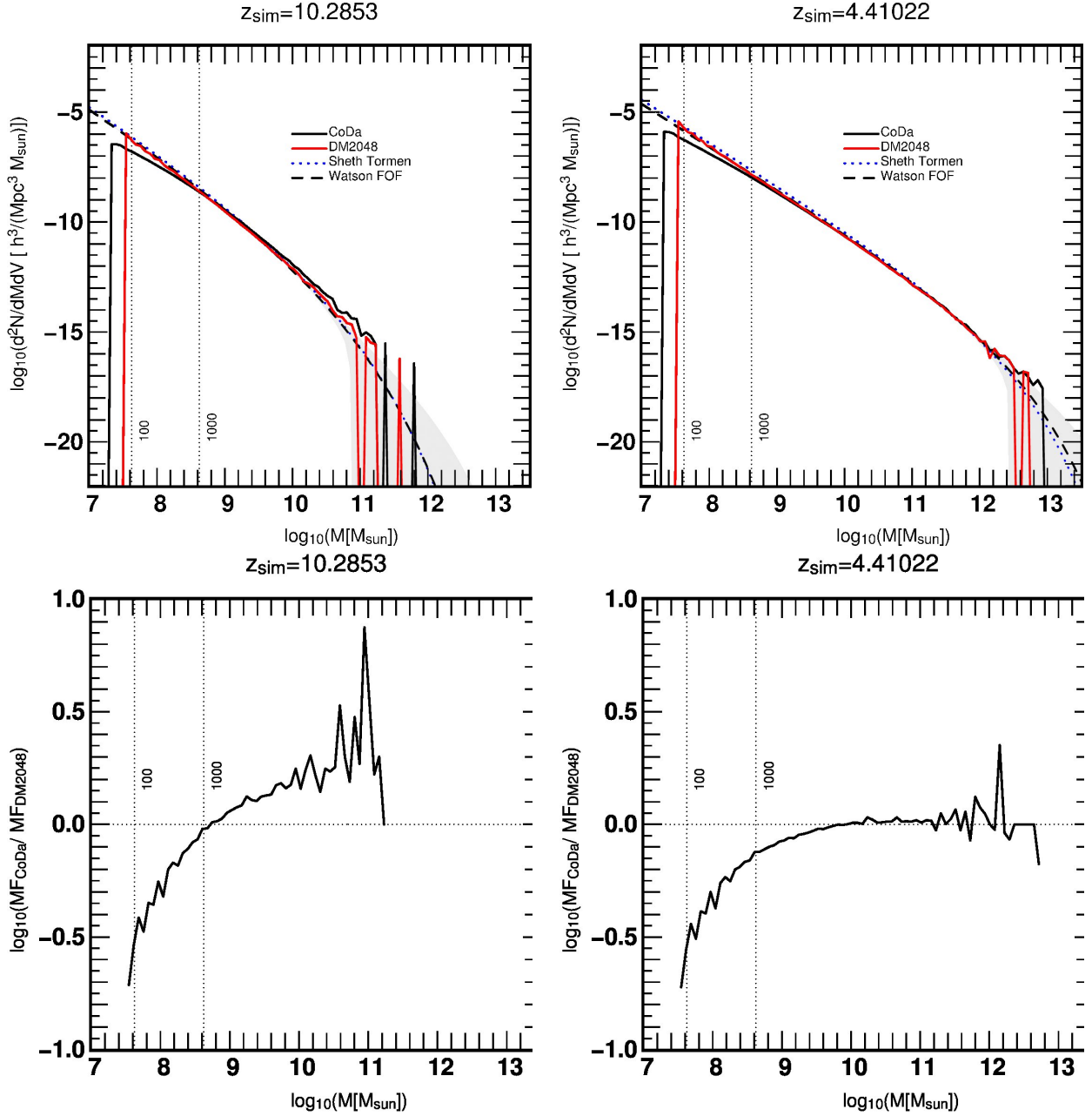


Figure A1. Mass functions of our simulations and comparison to literature, at $z_{\text{sim}}=10.28$ and $z_{\text{sim}}=4.41$. *Top:* the black and red solid lines show the mass functions obtained for CoDa and DM2048 respectively. Over-plotted are the theoretical mass function of [Sheth & Tormen \(2002\)](#) and the [Watson et al. \(2013\)](#) FoF universal fit (dotted and dashed lines, respectively). The gray area shows the Poisson error bars expected for the Watson FoF mass function. The vertical lines show the mass corresponding to 100 and 1000 particles haloes in CoDa. *Bottom:* The ratio of the the CoDa and the DM2048 mass functions is shown (black solid line).

sponds to a fixed interval of time just after this sharp end of reionization. Otherwise, if we compared them, instead, at the same cosmic time, haloes in different cases would have spent a very different amount of time experiencing the feedback effects of reionization, exposed to the UV background radiation. For instance, the end of reionization in the $32 \text{ h}^{-1}\text{Mpc}$ simulation happens at $z_{\text{sim}} \sim 4.6$, so at $z_{\text{sim}} = 4.5$, haloes have seen the post-reionization UV background for about 35 Myr. By contrast, for the highest resolution simulation, in the $4 \text{ h}^{-1}\text{Mpc}$ box, reionization ends at $z_{\text{sim}} \sim 5.8$, so at

$z_{\text{sim}} = 4.5$, haloes would have seen the post-reionization UV background for about 320 Myr, i.e. almost ten times longer. In order to make a meaningful comparison, then, we pick the time of comparison to be at the same interval of time just after each case's reionization ends. The redshifts chosen to compare halo properties for different cases are shown by the dot on each simulation's neutral fraction evolution (Fig. B1, left panel) and listed on the plot in the right panel, of the average SFRs as a function of halo mass for each case at its corresponding redshift.

The higher star formation rates for cases with higher resolution also mean that the vertical scale of the SFRs in Fig. B1 should be adjusted to make a direct comparison of the mass-dependence of the SFRs for different cases. This, too, is straightforward, since all of the cases have in common a universal shape for this mass dependence, which shares the slope of the SFR at high mass, above the mass scale where suppression occurs, and a turn-over at low-mass where suppression is occurring. This is obvious in Fig. B2, where we have renormalized the curves in Fig. B1 (right panel) by adjusting their vertical heights so as to make their high-mass SFRs lie on top of the power-law fit to this high-mass end of the CoDa SFR, given by $\log_{10}(\text{SFR}) = 5/3 \log_{10}(M) - 18.6$, and shown as a straight line on the log-log plot in both figures. It is clear from Fig. B2, now, as we compare the cases from lowest resolution to highest resolution, that the turn-over, reflecting the suppressed mass range (where Fig. B2 labels the value of the mass at which the low-mass turn-over segment joins the high-mass segment in each case), moves to successively lower mass, until the CoDa resolution is reached in the $8 \text{ h}^{-1} \text{ Mpc}$ box. When the resolution is increased yet again so as to exceed that of the CoDa simulation, by a factor of two in length and a factor of eight in mass, in the $4 \text{ h}^{-1} \text{ Mpc}$ box, the SFR dependence on mass is identical with that for the CoDa resolution case, as the curves for both cases completely overlap at all halo masses. This demonstrates that the suppression mass found by the CoDa simulation is well-enough resolved and not an artifact of numerical resolution.

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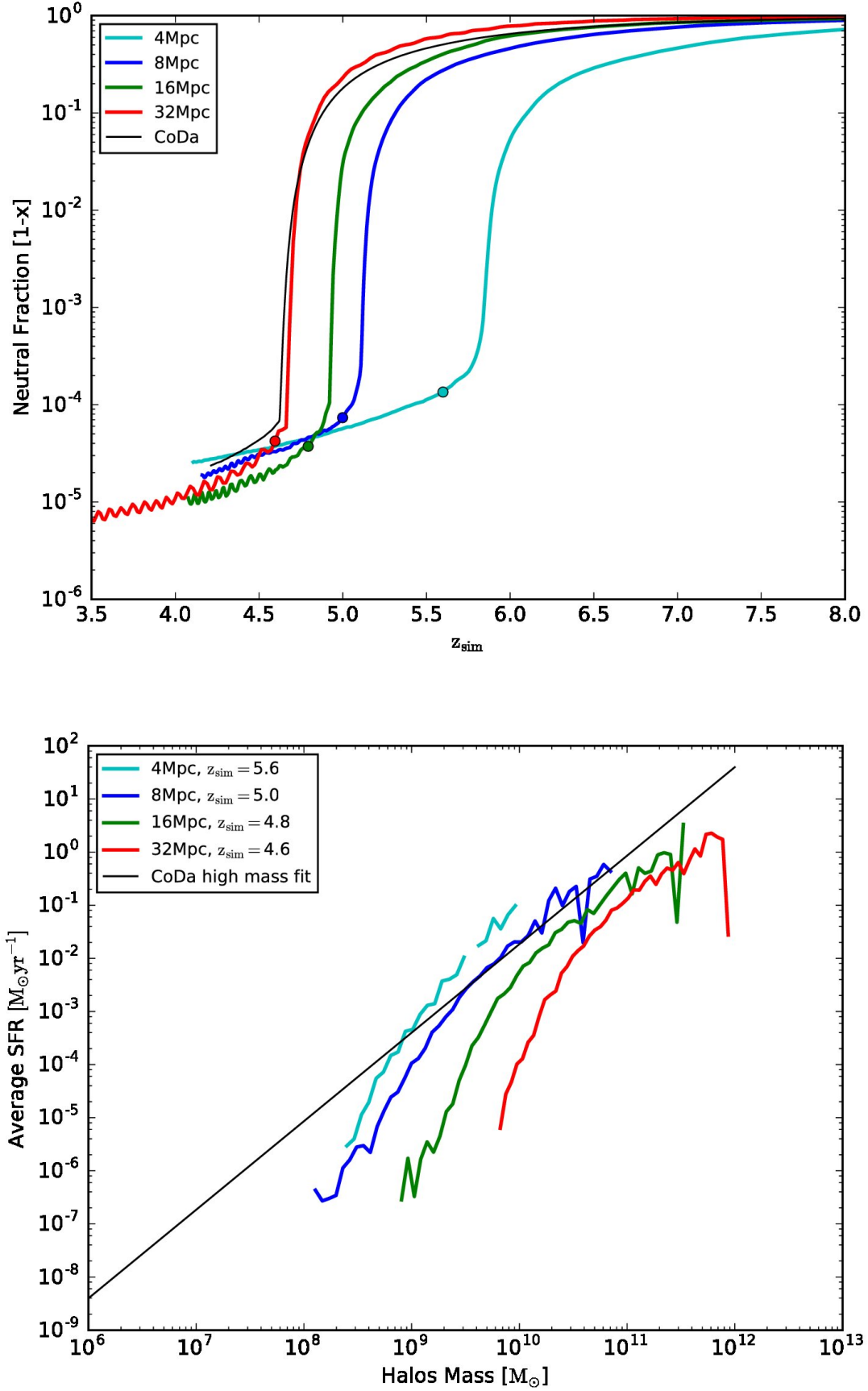


Figure B1. Neutral fraction and average SFR per halo for the resolution study.

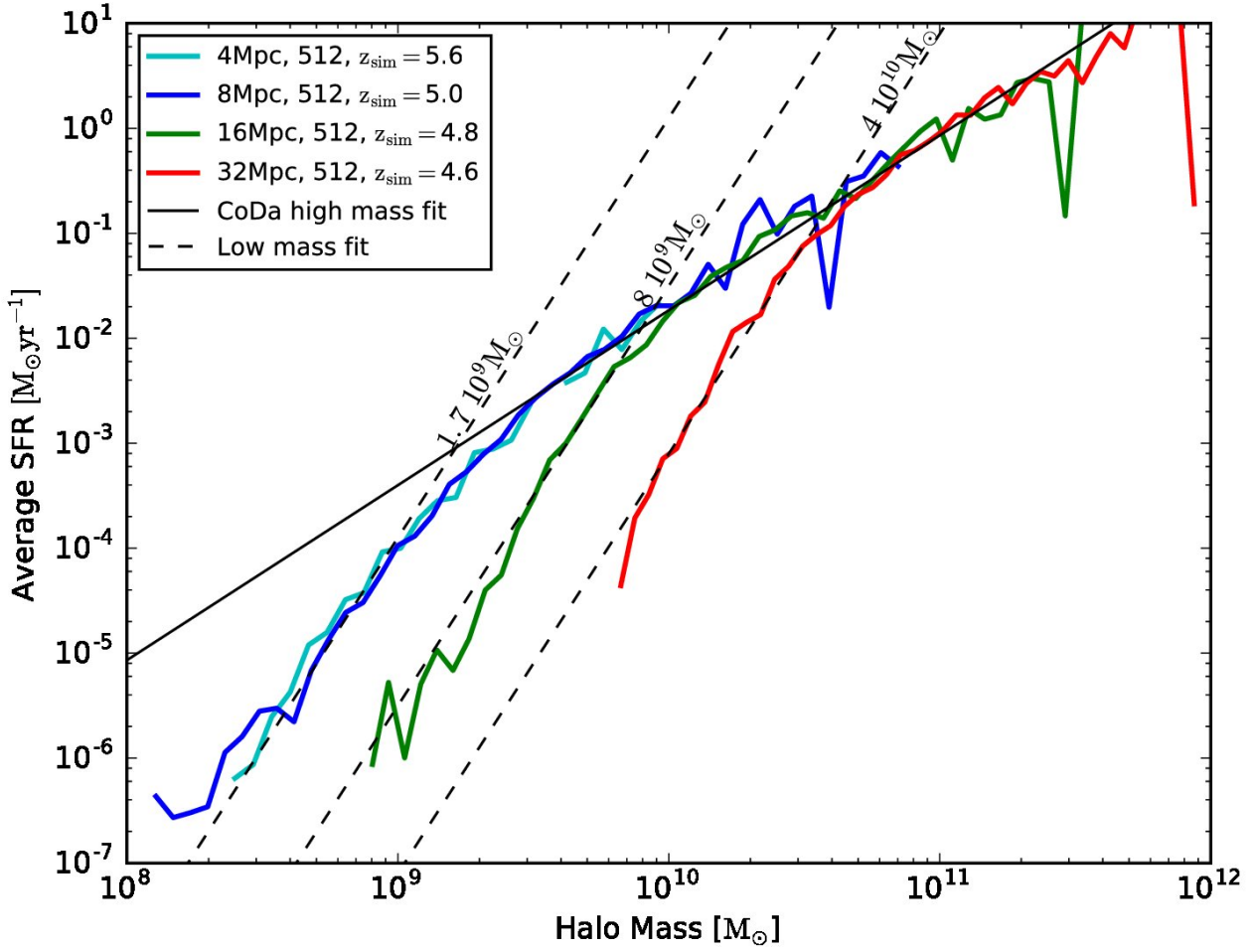


Figure B2. Average SFR per halo renormalized to match the high-mass branch of the $8 \text{ h}^{-1} \text{Mpc}$ simulation (CoDa resolution, solid black line, given by $\log_{10}(\text{SFR}) = 5/3 \log_{10}(M) - 18.6$). The dashed lines show a power law fit to the low mass branch of each simulation, with a slope M^4 . The numbers next to the dashed lines give the suppression mass as the mass where the low mass fit intersects the high mass fit.